Optical discovery and multiwavelength investigation of supernova remnant MCSNR J0512– 6707 in the Large Magellanic Cloud

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ABSTRACT

We present optical, radio and X-ray data that confirm a new supernova remnant (SNR) in the Large Magellanic Cloud (LMC) discovered using our deep Hα imagery. Optically, the new SNR has a somewhat filamentary morphology and a diameter of 56×64 arcsec (13.5×15.5 pc at the 49.9 kpc distance of the LMC). Spectroscopic follow-up of multiple regions show high [S II]/Hα emission-line ratios ranging from 0.66 ± 0.02 to 0.93 ± 0.01, all of which are typical of an SNR. We found radio counterparts for this object using our new Australia Telescope Compact Array (ATCA) 6 cm pointed observations as well as a number of available radio surveys at 8640 MHz, 4850 MHz, 1377 MHz and 843 MHz. With these combined data we provide a spectral index $\alpha \approx -0.5$ between 843 and 8640 MHz. Both spectral line analysis and the magnetic field strength, ranging from 124 to 184 μG, suggest a dynamical age between $\sim$2200 and $\sim$4700 yr. The SNR has a previously catalogued X-ray counterpart listed as HP 483 in the *ROSAT* Position Sensitive Proportional Counter (PSPC) catalogue.

Key words: surveys – ISM: individual objects: MCSNR J0512-6707; RP1577 – ISM: supernova remnants – Magellanic Clouds.

1 INTRODUCTION

Supernova remnant (SNRs) represent a prime source of heavy elements which are synthesized in their high-mass progenitors. They therefore play a major role in the enrichment of their host galaxies. Supernova (SN) explosions produce shocks, powerful blast waves and a radiation front strong enough for the resulting SNRs to shape and heat the interstellar medium (ISM), impacting on the ionization balance of the entire area in which they reside. SNRs usually produce radiation across the whole electromagnetic spectrum with detections in the optical, infrared (Lakićević et al. 2015), radio (due to non-thermal radiation e.g. Stupar et al. 2007), X-ray (detected by satellites e.g. Aschenbach 2002), the region of soft gamma-rays (due to the presence of residual young neutron stars) and even in the ultraviolet (Blair et al. 2006). Multiwavelength studies of SNRs provide information on several crucial aspects of stellar evolution and the physical behaviour of shocked material. They have a major influence on local magnetic fields whilst the shock waves may trigger new star formation and accelerate cosmic rays.

Within SNRs there is a correlation between Hα and radio-continuum emission even though the detailed fine structures do not overlap very often (see Cram, Green & Bock 1998). If the SN blasts into a hot, low-density ISM, Hα emission may be faint but if the ISM is cool and dense, Hα radiation is usually detected (with an intensity dependent on the density and temperature at the interface) revealing the shape and density of the shock front in the form of narrow filaments, elongated arc structures, bubble-like features and knots. These same structures can usually be seen in the forbidden line emission from [S II] 6717/6731 and to a more varying degree in [N II] 6548/6584, [O I] 6300/6364, [O III] 4959/5007 and the [O II] doublets close to 3727 Å and 7325 Å. Following Fesen, Blair & Kirshner (1985), a ratio of [S II] to Hα > 0.5 implies a shock front indicative of an SNR as does strong [O I] and [O II] lines plus the supporting presence of [N II] (Sabbadin, Minello & Bianchini 1977). In addition, the detection of Hα, rather than [O III], leading the radiation front would support the rejection of any Wolf–Rayet (WR) shell scenario (Gruendl et al. 2000).

At a distance of $\sim$49.9 kpc (Reid & Parker 2010 and references therein) the Large Magellanic Cloud (LMC) is an ideal laboratory to study the different varieties of SNRs in great detail across an entire galaxy. The LMC is close enough that SNRs are resolved sufficiently for multiwavelength studies to be conducted. Since SNRs

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yield key information on the death rate of high-mass stars and the number of white dwarfs exploding as Type Ia SN, an improved inventory of LMC SNRs, complete as possible, is required to put constraints on these values and improve our understanding of the final stages of stellar evolution. The LMC SNR sample is currently estimated at 59 remnants.[1]

The deep, stacked UKST Hα survey used to search for planetary nebulae in the LMC (Reid & Parker 2006a, Reid & Parker 2006b, 2013), also revealed several candidate SNRs, the first of which we report here, thanks to the optical depth of the Hα survey to R_{equiv} mag 22, and the method of image analysis. Spectroscopic followup of these objects (originally identified as SNR candidates in 2004) using 2dF on the Anglo-Australian Telescope (AAT) and the 1.9 m telescope at the South African Astronomical Observatory (SAAO) have permitted us to confirm identifications and conduct an optical spectroscopic and multiwavelength analysis.

The rest of the paper is structured as follows. In Section 2 we present the optical observational data including images and spectra. In sub-section 2.1 we derive an estimate for the expansion velocity and age of the SNR and in 2.2 estimate the radial velocity from key emission lines. In Section 3 we provide supporting radio data, the results of infrared imaging and previous X-ray observations. Section 4 is our conclusion where we summarize the combined multiwavelength data used to achieve our classification.

2 OPTICAL OBSERVATIONS AND DISCOVERY

The new SNR candidate, given the Reid & Parker reference name RP1577 and presented here for the first time, was discovered using deep-stack Hα and short red (SR) images of the LMC, SMC and surrounding regions (Reid & Parker 2006a). An Hα interference filter of exceptional specification and quality (Parker & Bland-Hawthorn 1998) was used to cover a substantial fraction of the UKST's large field. Its central wavelength (6590 Å) and bandpass (70 Å FWHM) work effectively in the UKST's fast f/2.48 converging beam. With a peak transmission of >85 per cent, the Hα filter is also partially sensitive to [N II]6583 Å. Overlapping fields with non-standard 4 deg centres enabled full contiguous coverage in Hα despite the circular aperture. High-resolution, panchromatic Tech-Pan film with peak sensitivity at Hα was used as the photographic detector (see Parker et al. 2005, for more details).

Over a period of three years, from 1997, a series of 12 repeated narrow-band Hα and 6 matching broad-band SR exposures of the central LMC field were also taken (see Reid & Parker 2006a). The purpose was that these images should be stacked to produce a single Hα and SR image with an overall depth increase of 1 magnitude over a single image frame. The twelve highest quality and well-matched UK Schmidt Telescope 2 h Hα exposures and six 15 min equivalent *SR*-band exposures were selected. From these exposures, deep, homogeneous, narrow-band Hα and matching broad-band SR maps of the entire central 25 deg^2 of the LMC were constructed.

The 'SuperCOSMOS' plate-measuring machine at the Royal Observatory Edinburgh (Hambly et al. 2001) scanned, co-added and pixel matched these exposures creating 10 μm (0.67 arcsec) pixel data which goes 1.35 and 1 magnitudes deeper than
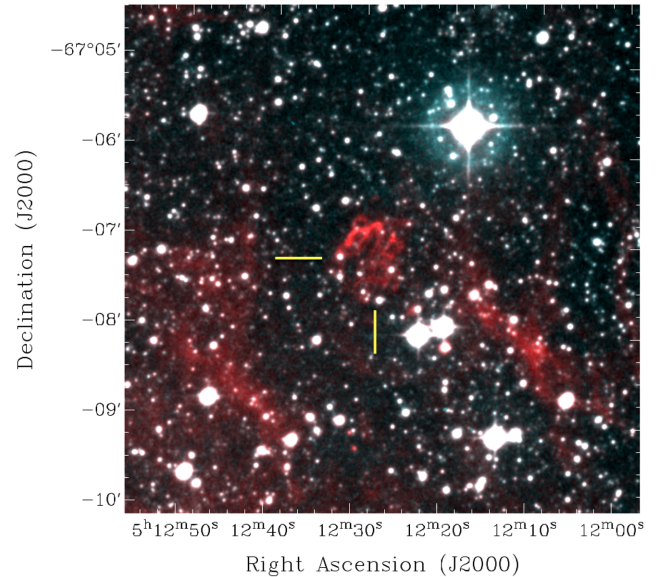
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[1] Estimated by the authors using 31 SNRs from *Chandra* catalog (hea-www.cfa.harvard.edu), 54 from Badenes, Maoz & Draine (2010) and recent discoveries including those from Maggi et al. (2014), Bozzetto et al. (2014a, 2014b), Warth et al. (2014) and Kavanagh et al. (2015).

Figure 1. The median stacked and combined Hα and broad-band red optical image of RP1577 where Hα is shown as red and the short red is shown as white.

individual exposures, achieving the full canonical Poissonian depth gain, e.g. Bland-Hawthorn, Shopbell & Malin (1993). This gives a depth ∼21.5 for the SR images and R_{equiv} ∼22 for Hα (4.5×10^{-17} erg cm^{-2} s^{-1} Å^{-1}) which is at least 1 magnitude deeper than the best wide-field narrow-band LMC images previously available. An accurate world coordinate system was applied to yield sub-arcsec astrometry, essential for success of the spectroscopic follow-up observations.

The resulting Hα and SR maps were originally overlayed with false colours in order to detect planetary nebulae (Reid & Parker 2006a, Reid & Parker 2006b, 2013) but also proved an excellent tool for uncovering a large number of hot and cool emission-line stars (Reid & Parker 2012). In addition to these discrete objects, the maps were able to reveal the fine morphological characteristics of SNRs, which appear as complex interwoven streams, elongated arcs, bubbles and narrow filaments in Hα. The newly identified candidate is shown in Fig. 1 where Hα is coloured red and the short red continuum image appears white due to both blue and green colours being assigned to the same image. Although these structures can be closely correlated with radio emission, they usually become quite fragmented in old ;remnants and may extend beyond, within or along the boundaries of the radio remnant (see Stupar & Parker 2009). In addition to the Hα/SR combined images, we used the Magellanic Cloud Emission-Line Survey (MCELS; Smith et al. 1998) to directly compare the strength of the three narrow bands of [O III] 5007 (blue), Hα (red) and [S II] 6716+6731 (green). When the colour maps are evenly combined the SNR candidate becomes evident by the strong green ([S II] line) colour.

RP1577 has a central position of (J2000) RA 05^{h}12^{m}28^s DEC −67°07 arcmin 20 arcsec determined using SuperCOSMOS astrometry (Hambly et al 2001) and a measured angular extent from the Hα imagery of 56 × 64 arcsec. This equates to a physical extent of 13.5 × 15.5 pc at the LMC distance. With physical diameters ranging from 0 to 120 pc (Badenes et al. 2010), the cumulative distribution of SNR diameters in the LMC have been found to be roughly linear (eg. Chu & Kennicutt 1988; Bandiera & Petruk 2010) up to a cutoff at a physical radius of r_{cut} ∼30 pc (Badenes et al.

Table 1. Observing Logs for the follow-up of emission object and SNR candidate RP1577 in the LMC. The first column gives the telescope name, second gives the instrument name, third gives the observation date, the fourth gives the grating name, the fifth gives the dispersion in angstroms, the sixth gives the central wavelength, the seventh gives the wavelength coverage in angstroms, the eighth gives the exposure time in seconds, the ninth gives the number of exposures taken.

Telesc.	Inst.	Date	Grating dispenser	Dispersion Å/pixel	Central λ (Å)	Coverage λ (Å)	T_{exp} s	N_{exp}
AAT	2dF	04-Dec 16	300B	4.3	5852	3660–8000	1200	3
AAT	2dF	04-Dec 18	1200R	1.1	6793	6220–7340	1200	2
1.9-m	ccd	14-March 7	300	5	5800	3850–7738	800	2

Table 2. Observed line intensities and flux ratios for four different spectroscopically observed positions across three areas of RP1577. Line measurements were made using the SPLOT script in the IRAF package. Errors for individual flux line measurements are a constant $\pm 1.9 \times 10^{-17}$ for 2dF 300B, $\pm 1.2 \times 10^{-17}$ for 2dF 1200R, $\pm 1.5 \times 10^{-17}$ for AAOmega, $\pm 2 \times 10^{-17}$ for 1.9 m measurements. Column details are provided in the text.

Obs.	RA (J2000)	DEC (J2000)	Hβ	[O III]$_{5007}$	Hα	[N II]$_{6583}$	[N II]/ Hα	[S II]/ Hα	[S II] ratio 6716/6731	n_e (cm^3)	T_e (K)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)
2dF–300B	05 12 28.4	−67 07 29	2.22E–15	3.84E–15	6.95E–15	1.49E–15	0.32 ± 0.02	0.46 ± 0.02	1.35 ± 0.03	70 ± 9	35 000 ± 10 000
2dF–1200R	05 12 28.4	−67 07 29	–	–	–	–	0.25 ± 0.02	0.66 ± 0.02	1.44 ± 0.03	<26	–
1.9 m 1	05 12 31.0	−67 06 59	1.81E–14	2.20E–14	6.40E–14	1.16E–14	0.24 ± 0.01	0.78 ± 0.01	1.40 ± 0.02	26.5 ± 8	–
1.9 m 2	05 12 26.0	−67 07 05	3.2E–15	7.17E–15	1.17E–14	2.32E–15	0.30 ± 0.01	0.93 ± 0.01	1.44 ± 0.02	<26	–

2010). This means that the 15.5 pc size of RP1577 places it among the smaller 15 per cent of LMC SNRs, joining six other SNRs in the 10–20 pc range.

To eliminate the possibility that RP1577 was the product of a WR shell we checked the available data bases for the presence of any WR star or candidate inside (or in the vicinity of) the shell but none were found within a three arcmin radius. We then compared the expanded radius of Hα and [O III] emission since, according to theoretical models, the outermost [O III] usually leads the Hα emission in WR shells (Gruendl et al. 2000). This occurs due to a decrease in the sensitivity of Hα at temperatures greater than 10^4 K while [O III] emission is unaffected by the higher temperatures. As a result, the morphologies of SNRs and WR shells should be somewhat different (see for example Cox 1972, Gruendl et al. 2000). As these models depend on the interaction of the shells with the ISM, they remain as a useful diagnostic whether or not they are used to identify WR shells. Using the MCELS survey we overlayed the [O III] and Hα maps in order to check whether [O III] was leading at any point around RP1577. Despite the low 2.5 arcsec/pixel resolution of the MCELS survey, a pixel by pixel comparison along the shock front showed that Hα was in fact leading the [O III] emission, clearly supporting the rejection of any WR shell scenario.

In Table 1 we summarize details regarding the spectroscopic follow-up observations. A five night observing run on the AAT using 2dF (Lewis et al. 2002) was undertaken in 2004 December to spectroscopically confirm LMC emission candidates. The large corrector lens incorporates an atmospheric dispersion compensator, which is essential for wide wavelength coverage using small diameter fibres. Specifically targeted long-slit spectra were obtained using the 1.9 m telescope at the SAAO in 2014 March.

Optical properties for the new SNR in the LMC are given in Table 2. Column (1) indicates the telescope and observing method used at each position. The AAT observations were conducted using the 2dF multifibre spectroscopy system where each fibre covered 2 arcsecs diameter on the sky. Column (2) gives the (J2000) position of the optical centre of the observed position in right ascension. Column (3) likewise, gives the (J2000) declination of the same position. Each observed position was determined using the Hα map and

checked against SuperCOSMOS astrometry. A comparison of the resulting spectral signature was intended to provide a good overall analysis of the ratio of [S II] to Hα and [N II] across the nebula, all measured with Gaussian line fitting from within the IRAF SPLOT script. Column (4) gives the measured Hβ flux at the position of each slit or fibre. Column (5) gives the measured flux of the [O III] 5007 emission line at the observed position. Flux calibration of the 2dF line intensity was conducted using the method outlined in Reid & Parker (2010). For the purpose of diagnostics the Hα flux is provided in column (6), followed by the [N II] flux in column (7). Column (8) then gives the [N II] 6583+6548/Hα ratio including the error based on line measurements. It should be noted that the [N II] 6548 line is quantum mechanically fixed at a 1–3 intensity ratio with the [N II] 6583 line. Column (9) provides the [S II]/Hα ratio and line measurement errors for estimating the excitation of the shock-front. Column (10) shows the ratio of [S II] 6716/6731 used for estimating the electron density. The electron density shown in column (11) is derived from the [S II] 6716/6731 ratio derived using the TEMDEN script found in the STSDAS package from IRAF. Once the [S II] ratio reaches 1.4, representing a density of 26 cm^3 the low-density limit using this method has been reached. Column (13) gives the electron temperature, derived from the [O III] 5007+4959/[O II] 4363 emission lines. The faint [O II] 4363 line was only detectable for the purpose of measurement at one position.

The optical Hα/Hβ ratio at each observed position was used to determine the extinction constant c(Hβ) (i.e. the logarithmic extinction at Hβ) at that position. These hydrogen transitions are the strongest and easiest to accurately measure in the nebula spectrum and are fortunately located very close to the other main emission lines used for diagnostics. The observed Hα/Hβ ratio, when compared to the recombination value of 2.86 (Aller 1984), gives a logarithmic extinction at Hβ of:

$$c(H\beta) = (log(H\alpha/H\beta) - log(2.86))/0.34. \quad (1)$$

This estimation is based on the relationship between observed and intrinsic intensities:

$$\frac{I_{obs}(H\alpha)}{I_{obs}(H\beta)} = \frac{I_{int}(H\alpha)}{I_{int}(H\beta)} 10^{-c(H\beta)[f(H\alpha)-f(H\beta)]}, \quad (2)$$
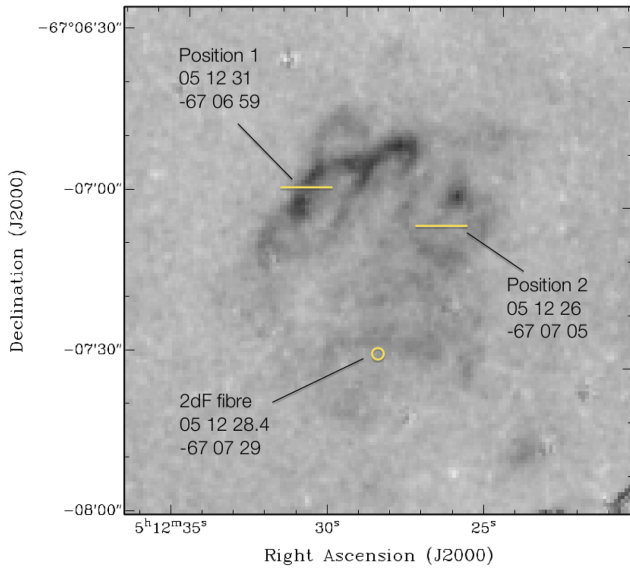
Figure 2. The (Hα/SR) quotient image of RP1577 showing circular loops to the north of parallel, stratified filaments which outline the exterior of the expanding shell. The position of the 2dF spectroscopic observation is shown by the yellow circle close to the centre of the object. Horizontal lines mark the positions of long-slit observations.
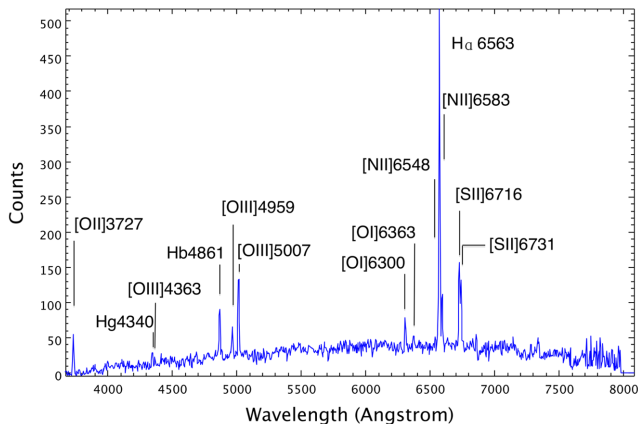

Figure 3. The 2dF spectroscopic observation of RP1577 indicated by the yellow circle in Fig. 2. The Hα/[S II] ratio is 0.46 ± 0.02. The extinction of $c(H\beta) = 0.45$ is based on the Balmer decrement.

where $[f(H\alpha)\text{-}f(H\beta)] = -0.34$ from the standard interstellar extinction curve. The resulting extinction values $c(H\beta)$ are given in the caption below each spectrum in Figs 3, 5 and 6.

The [S II] 6717/6731 ratio of 1.35 ± 0.07 from the 2dF 300B exposures results in an electron density value near 70 ± 10 cm^3. In Fig. 2 we show the (Hα/SR) quotient image which effectively removes stars from the area. The yellow circle in the image marks the position observed with 2dF. This is approximately 10 arcsec south of the optical centre as indicated by the elliptical shape of the diffuse Hα emission. Spectroscopic results from this position are shown in Fig. 3 for the 300B exposure and Fig. 4 for the 1200R exposure.

Both 2dF spectra show characteristics and ratios that are quite typical of many other SNRs we have observed within the LMC. The [N II]/Hα and [S II]6716/6731 line ratios from Table 2 are in full agreement with those of other LMC SNRs (Payne, White & Filipović 2008). At the position observed by 2dF the [S II]/Hα ratio
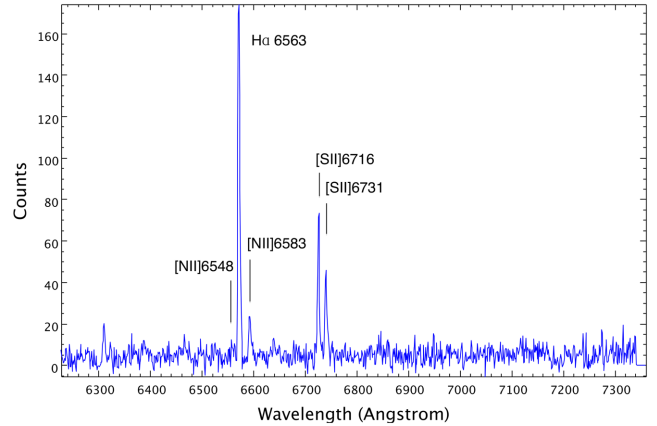

Figure 4. The 2dF 1200R medium resolution spectroscopic observation of RP1577 indicated by the yellow circle in Fig. 2. The higher resolution of this spectrum shows the weak [O I]6300 & 6363 lines. The [S II]/Hα ratio is 0.66 ± 0.02.

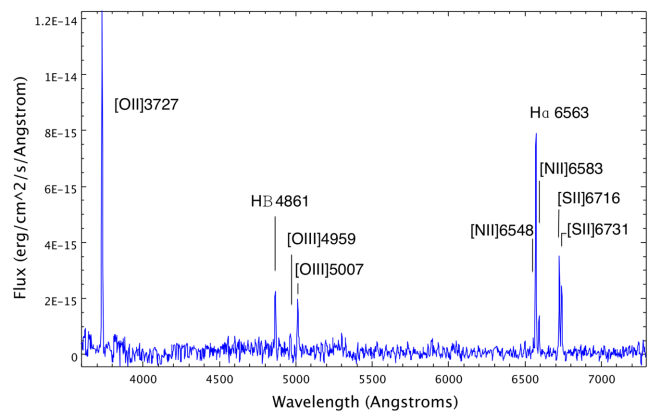**Figure 5.** The long-slit, spectroscopic observation of RP1577 at position (J2000) RA $05^h12^m31^s$ and DEC $-67°06'59''$. The [S II]/Hα observed ratio at this position is 0.78 ± 0.01. The extinction of $c(H\beta) = 0.28$ is based on the Balmer decrement.

of 0.46 ± 0.02 from the 300B grating is commonly expected for an SNR (Fesen et al. 1985). Higher [S II]/Hα ratios of 0.78 and 0.93 were obtained from the 1.9 m spectra in the northern regions of the object. In these positions our image data show dense filaments of emission which can be created in an SNR by post-shock plasma. Just below the centre of the SNR, the shocks have a different effect, increasing the electron temperature and ionization structure of the recombination region, as also seen in the more prominent [O I] 6300 lines in Figs 3 and 4. At the adiabatic stage in the evolution of the SNR we would expect the photoionization fields to be rather weak close to the centre where photons are re-ionizing gas which is seeping back into the shock. This probably explains why the [O I] 6300 line is higher in the central region but lower in the outer shells where the [S II]/Hα ratios appear to be higher. The [O I] 6300 lines in the raw 1.9 m long-slit spectra are wide near the continuum therefore their low level in the reduced spectra (Figs 5 and 6) may also in part be due to sky subtraction where the combined nebulae and telluric [O I] 6300 lines span the width of the slit. Morphologically, although the overall shape is somewhat elliptical, the interior detail of the SNR shows a series of elongated filaments, which become more dense and finally curved in the form of loops to the north. With such strong spectroscopic confirmation, we now assign RP1577 the
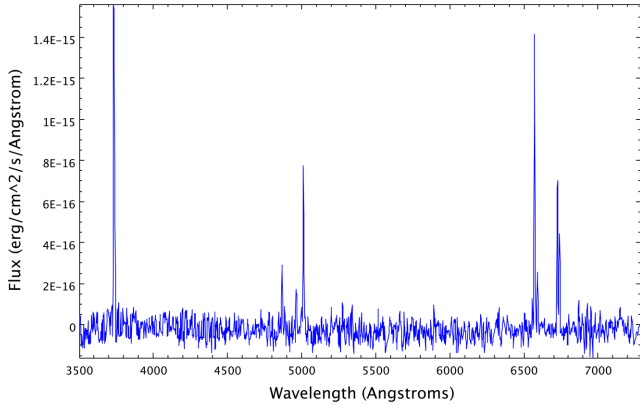
Figure 6. The long-slit, spectroscopic observation of RP1577 at position (J2000) RA 05^{h}12^{m}26^s and DEC $-$67°07′05″. The [S II]/Hα observed ratio of 0.93 $\pm$ 0.01 at this position is the highest ratio of all our spectroscopic observations of RP1577. The extinction of c(Hβ) = 0.33 at this position is based on the Balmer decrement.

name MCSNR J0512−6707which includes the basic coordinates as recommended by the IAU.

2.1 Expansion velocity and age

To estimate the expansion velocity and the dynamical age, we first assume that the SNR is currently in the Sedov–Taylor phase (Sedov 1959) where the swept-up mass is large compared to the ejected mass, the shock velocity is decelerating and the evolution is approximately adiabatic. The dynamical age is then $t = C \times R / V_s$, with C as a constant which depends on the expanding velocity of shock wave after the SN explosion. This method, often used for shell-type remnants in the adiabatic stage would require C to be a value of 2/5. In this scenario MCSNR J0512−6707 is still in the Sedov–Taylor phase, which is typical for SNRs up to ∼30 000 yr of age and diameters up to ∼30 pc (Badenes et al. 2010). As such the radii should expand as

$$r \sim E_0^{1/5}\rho^{-1/5}t^{2/5}, \tag{3}$$

where E_0 is the kinetic energy of the original SN, ρ is the ambient gas density and t is time. The shock velocity decelerates according to

$$V_s = \frac{dr}{dt} \sim E_0^{1/5}\rho^{-1/5}t^{-3/5}. \tag{4}$$

Our adopted ambient initial gas density is very difficult to estimate given the location of the SNR very close to a dense H II region and surrounded by dense dust emission (see Section 3.3). Estimates may be complicated by inhomogeneous regions of nebula with varying optical thickness or optically thick cloudlets of dust. For a close approximation to the average initial H density, we adopt the 1 cm^{-3} ambient gas density found for use with LMC SNRs by Berkhuijsen (1987). This value may be compared to the initial H density value of 10 cm^{-3} for H II regions in the LMC as found by Pellegrini et al. (2011) which takes into account gas-phase abundances and the gas-to dust ratio averaged across the densest regions. To estimate an error in the ambient gas density we allow for a rise of up to 1.2 cm^{-3} due to dense dust in the immediate area. This was estimated by measuring diffuse emission in the general 30 arcmin radius surrounding MCSNR J0512−6707 using our 2dF fibres. The [S II] 6716, 6731lines provide electron densities (n_e) which are proportional to the square root of the emission-line vol-

Table 3. Heliocentric radial velocities and systematic errors for key emission lines at the observed positions as shown in Fig. 2.

Obs.	[O III] (5006.84Å) (km s^{-1})	Hα (6562.80Å) (km s^{-1})	[N II] (6583.41Å) (km s^{-1})	[S II] (6716.47Å) (km s^{-1})	error (km s^{-1})
2dF-1200R	–	377	391	396	$\pm$ 12
Posn. 1	332	341	376	365	$\pm$ 28
Posn. 2	312	390	407	411	$\pm$ 28

ume emissivity. These measurements provide a clue as to whether the ambient environment of MCSNR J0512−6707 is proportionally more dense than average. The electron density n_e of ambient H II is between 55 cm^{-3} and the low density limit of = 26.0 cm^{-3} which is slightly higher than the standard LMC measurement of faint, diffuse emission (< 26.0 cm^{-3}). Since we find that diffuse H II regions have a low density n_e = ∼600 cm^{-3}, and a H density of 7 cm^{-3} (see above), an $n_e \sim 38$ cm^{-3} on the same scale would be equivalent to a density of ∼1.2 cm^{-3}. For the lower error estimate we adopt a value of 0.2 cm^{-3} since a value of 0.3 cm^{-3} was found by plotting the SNR on the correlation provided by Berezhko & Völk (2004).

The upper error for our adopted energy is estimated from the radio observations for this remnant (see Section 3) and based on a surface brightness of 1.57×10^{-20} W m^2 Hz^{-1} sr^{-1} at a frequency of 1 GHz, placing the object at a position on the $\Sigma - D$ diagram by Berezhko & Völk (2004) equivalent to an energy (E$_0$) of ∼0.25 × 10^{51} erg. We adopt this estimated value as it represents an upper estimate for such an SNR in the LMC. For our lower error we adopt a value of 9.5× 10^{51} erg, which is just below the canonical 10^{51} erg value for SNRs. The radius of 7.75 pc is found from the maximum diameter given above. This method provides an estimated dynamical age of the remnant between ∼2000 yr and ∼5000 yr. The method assumes that the swept-up mass will become proportionally larger than the ejected mass and the velocity will decelerate accordingly. Using the same equations, the expansion velocity may lie between ∼ 600 and 1400 km s^{-1}.

2.2 Radial velocity

The heliocentric radial velocities at each observed position of the SNR are shown in Table 3. Emission lines were measured using the *IRAF* splot and EMSAO tasks against standard rest velocities. Column 1 gives the name of the observation, which corresponds to the positions marked in Fig. 2. Columns 2 to 5 then present the heliocentric radial velocities in km s^{-1} for the [O III] 5007, Hα, [N II] 6583 and [S II] 6716 emission lines. The estimated velocity error, given in column 6 is a standard systematic error for all velocities measured on a particular observation. It was determined according to the size of the pixels, related to the resolution of the individual spectrum. The rest velocities for each element are provided in the column headers.

As measured through optical imagery in Section 2, above, the [O III] emission is less affected by temperatures greater than 10^{4}K and so, can be seen to move at a slower velocity than Hα which leads the shock front. This effect can be measured, even though expansion at positions 1 & 2 are directionally near to the plane of the sky. Our velocity measurements show that this is true of the [O III] line at both positions observed by the SAAO 1.9 m telescope. Velocities on the forbidden lines of [N II] and [S II] are in every case higher than those found for Hα. This is to be expected for
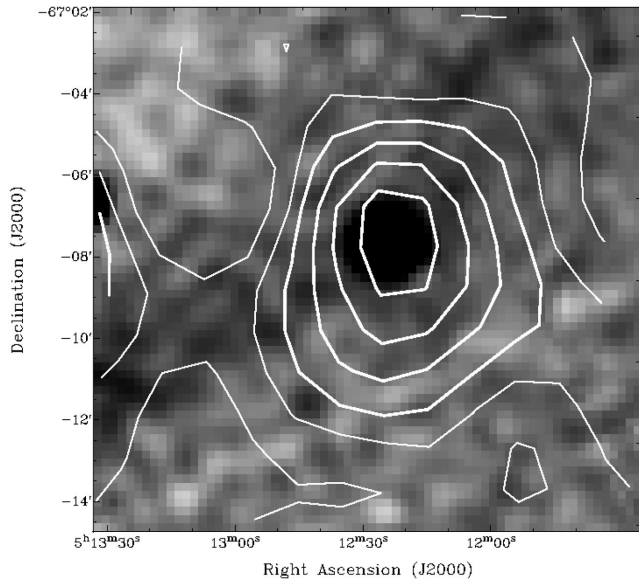
Figure 7. Radio contours from PMN at 4850 MHz superimposed on the 843 MHz SUMSS grey-scale map of SNR MCSNR J0512−6707. Contours are spaced at 0.01 mJy beam^{-1}. The flux density varies from −0.66 mJy to 3.16 mJy across the area shown.

[S II] since the bulk abundance of this element is produced in the SN explosion.

3 RADIO, INFRARED AND X-RAY OBSERVATIONS

3.1 Radio observations

Radio detection has traditionally been the main method for identification and confirmation of SNRs because of their non-thermal spectra, radio shell structures and polarization (e.g. Bozzetto et al. 2014b). Radio observations also alleviate problems associated with optical detection where dust often obscures optical detail. With the development of new instrumentation, the sensitivity of existing radio telescopes is constantly being enhanced, allowing us to improve our knowledge of SNRs within the LMC. For this project we have used a combination of dedicated radio observations complimented by freely available radio survey data.

The radio counterpart to this optically detected SNR was originally found using data from the Parkes-MIT-NRAO (PMN; beam size = 300″) radio survey at 4 850 MHz (Condon, Griffith & Wright 1993) and the Sydney University Molonglo Sky Survey (SUMSS; beam size = 45″) at 843 MHz (Bock, Large & Sadler 1999). Although a radio source is detected at the position of MCSNR J0512−6707 in both of these surveys, it is more extensive in the PMN due to the lower resolution of the survey (see Fig. 7). The contours therefore do not represent the true size. A detection of MCSNR J0512−6707 was made in Filipović et al. (1995; listed as LMC B0512−6710) where a radio continuum analysis of the LMC was performed at various frequencies using the Parkes radio telescope. The SUMMS 843 MHz measurement was omitted from spectral index calculations as it was superseded by the 843 MHz *Molonglo Sythesis Telescope* (*MOST*) measurement, due to the higher sensitivity of the MOST image and its greater UV–coverage.

The PMN and SUMSS radio archival data, though helpful in confirmation was of insufficient quality to reveal the structure and so the
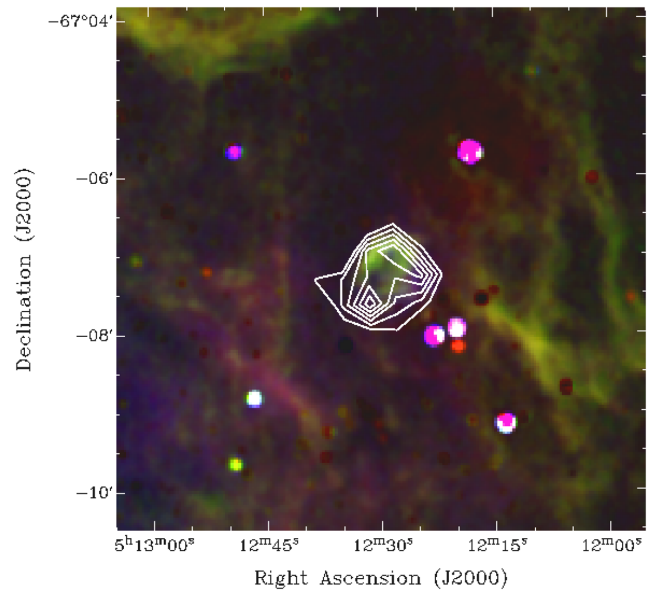

Figure 8. MCSNR J0512−6707 with contours from the Dickel et al. (2010) 8640 MHz mosaic image (beam size = 22 × 22 arcsec). The local rms is $\sigma = 0.5$ mJy beam^{-1} and contours are 3, 4, 5, 6, 7 and 8 σ. The radio contours are superimposed on the continuum–subtracted MCELS map of the immediate area where Hα is red, [O III] is blue and [S II] is green.

latest radio-continuum data and our own pointed observations were used to provide details and additional frequencies to demonstrate the non-thermal nature of MCSNR J0512−6707. Radio-continuum data used in this project includes a 36 cm (843 MHz) *MOST* mosaic image (as described in Mills et al. 1984), a 20 cm (1377 MHz) mosaic image from Hughes et al. (2007), and 6 cm (4 800 MHz) and 3 cm (8 640 MHz) mosaic images from Dickel et al. (2005, 2010) which are shown as contours against the continuum–subtracted MCELS three-colour map (see Section 2) in Fig. 8.

These mosaic data were complemented by pointed observations of MCSNR J0512−6707 from our Australian Telescope Compact Array (ATCA) observations taken on 2011 November 15 and 16, using the Compact Array Broad-band Backend (CABB) receiver which has a bandwidth of 2 GHz. The compact array configuration EW367 was chosen in an effort to obtain as much emission from the shorter spacings at 6 and 3 cm ($\nu = 5\,500$ and $9\,000$ MHz, respectively) as possible. The resulting 6 cm contours are overlaid on the MCELS Hα map of MCSNR J0512−6707 in Fig. 9. Total integration time over the two days for this source amounted to ∼64 min. The MIRIAD[2] (Sault, Teuben & Wright 1995) and KARMA[3] (Gooch 1995) software packages were used for reduction and analysis. Images were formed using MIRIAD multifrequency synthesis (Sault & Wieringa 1994) and natural weighting. They were deconvolved with primary beam correction applied.

To estimate the radio spectrum of MCSNR J0512−6707, integrated flux density measurements were taken from the four mosaics at frequencies of 843, 1377, 4800, and 8640 MHz, in addition to our observations at 500 and 9000 MHz. However, as interferometers such as the ATCA are an arrangement of individual antennas, they lack the short spacings that are responsible for the large-scale emission (e.g. of objects such as SNRs), resulting in a loss in flux density which becomes more pronounced at higher frequencies. For
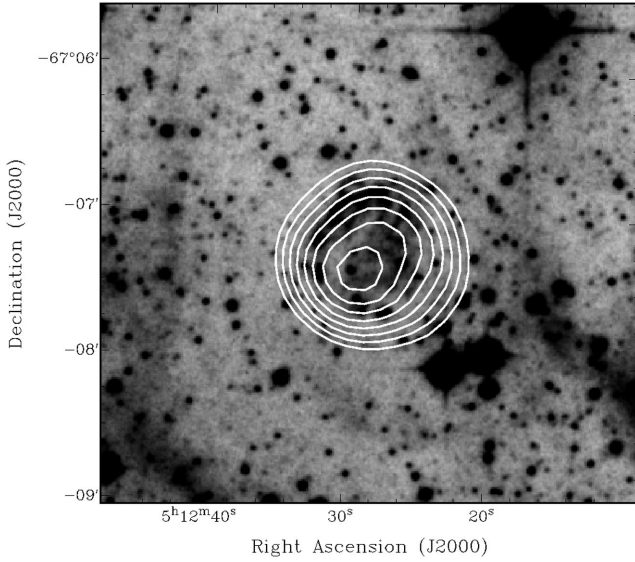
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[2] http://www.atnf.csiro.au/computing/software/miriad/
[3] http://www.atnf.csiro.au/computing/software/karma/

Figure 9. Contours using the entire 2 GHz bandwidth centred at 5 500 MHz are overlaid on the Hα image of MCSNR J0512−6707. The contours are spaced at ∼ 0.01 mJy beam^{-1} commencing at 0.3 mJy beam^{-1}.

Table 4. Integrated Flux Densities of MCSNR J0512−6707 used in this work.

ν (MHz)	λ (cm)	rms (mJy)	Flux density (mJy)	Beam Size (arcsec)	Image Ref.
843	36	0.8	115.5	46.4 × 43.0	MOST
1 377	20	1.1	86.9	40.0 × 40.0	Hughes et al. (2007)
4 732	6	0.2	49.6	44.0 × 28.4	This work
4 800	6	0.7	53.7	35.0 × 35.0	Dickel et al. (2010)
5 244	6	0.2	47.1	40.1 × 25.9	This work
5 756	6	0.2	44.3	36.6 × 23.7	This work
6 268	6	0.2	40.7	33.9 × 21.9	This work
8 640	3	0.7	29.6	22.0 × 22.0	Dickel et al. (2010)

this reason, the 9000 MHz integrated flux density measurements were discarded. The 5500 MHz observations were split into four 512 MHz bands (4732, 5244, 5756, and 6268 MHz), and integrated flux density measurements were taken per band with fitted gaussians using the MIRIAD task IMFIT. The results of these measurements and those taken from the mosaics are listed in Table 4 and shown in Fig. 10. Errors in these measurements predominantly arose from uncertainties in defining the 'edge' of the remnant. However, we estimate errors to be less than 10 per cent. The rms noise estimates were taken using the MIRIAD task SIGEST. The resulting spectral index between 843 and 8 640 MHz of $\alpha = -0.52 \pm 0.04$ is consistent with the typical SNR spectral index value of −0.5.

3.2 Magnetic field strength

We use the equipartition formula as given by Arbutina et al. (2012, and the corresponding 'calculator'[4]) to estimate the magnetic field strength of this SNR. However, further to being unable to accurately measure the extent of the remnant, the width of the shell was impossible to measure from a source which is somewhat unresolved in the radio, and consequently, the filling factor was unable to be estimated. Therefore, Table 5 shows the average equipartition
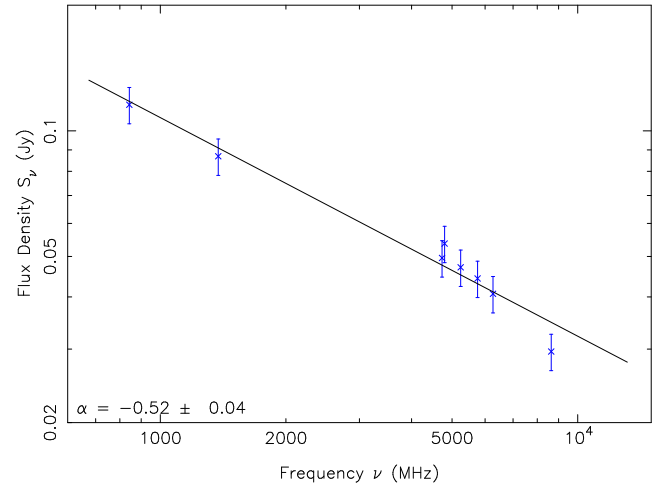
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[4] The calculator is available at http://poincare.matf.bg.ac.rs/~arbo/eqp/

Figure 10. Radio-continuum spectral index of MCSNR J0512−6707 between 843 and 8640 MHz. The blue markers represent flux density error margins of 10 per cent.

Table 5. Magnetic field strength and minimum energy for varied filling factors (*ff*).

ff	B (μG)	E_{min} (erg)
1.00	124	6.75×10^{49}
0.75	134	5.96×10^{49}
0.50	151	5.01×10^{49}
0.25	184	3.71×10^{49}

Table 6. Magnetic field strength for a sample of LMC SNRs with their associated ages.

SNR	Age (yr)	ΔAge (yr)	B (μG)	Reference
0509−6731	400*	120	168	Bozzetto (2014b)
0519−6902	600*	200	171	Bozzetto et al. (2012a)
0533−7202	17 500	3 500	45	Bozzetto et al. (2013)
0508−6902	22 500	2 500	28	Bozzetto (2014a)
0529−6653	37 500	12 500	48	Bozzetto et al. (2012b)

Note. *These age estimates are taken from Rest et al. (2005)

field over the whole shell of MCSNR J0512−6707 in addition to the estimated minimum energies based on various filling factors ranging from 1 (a completely filled SNR) to 0.25 (a mostly empty shell-like remnant). The range of values for the magnetic field, from 124-184 μG, are relatively high for an SNR, and may be indicative of a younger remnant. This can be compared with other LMC SNRs using the same method, e.g. those listed in Table 6. Such values may suggest a remnant of a few thousand years, which is equal to or less than the age estimate of ∼2280 to ∼4690 yr derived in Section 2.1 but still in-keeping with the remnant's relatively small size which ranks it as the 6th smallest currently known in the LMC (Badenes et al. 2010).

3.3 Dust

SNRs heat surrounding dust while mass maps indicate that they also destroy $3.7^{+7.5}_{-2.5}$ M$_\odot$ of dust per SNR through the process of sputtering (Lakićević et al. 2015). Type II SN are able to produce

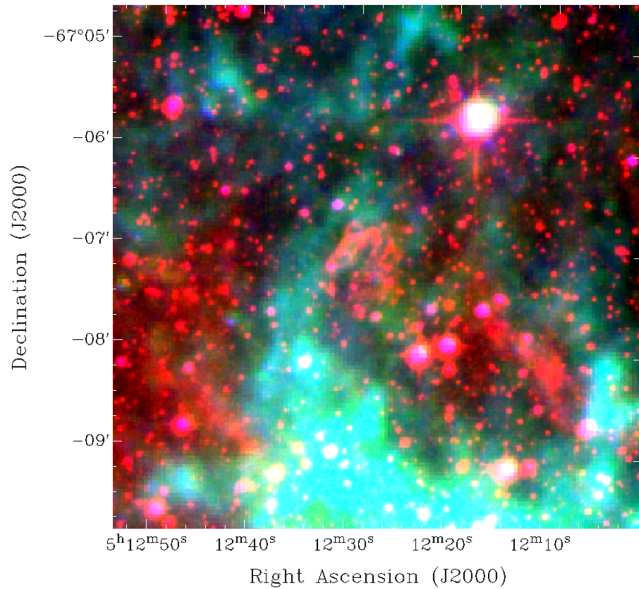**Figure 11.** The Hα map of MCSNR J0512−6707 and its surrounding area (shown in red) is overlaid with the 8 μm SAGE image (blue) and the 24 μm MIPS image (green).

dust on the order of a solar mass per explosion (Deneault, Clayton & Heger 2003; Nozawa et al. 2003). The average lifetime for interstellar dust under exposed to a SNe are estimated at $2^{+4.0}_{-1.3}$ years (Lakićević et al. 2015) indicating that the amount of dust that survives may be extremely low. Since SNRs have such a major influence on the local dust environment and effect a galaxy's evolution (Lakićević et al. 2015) it is worth examining the environment of MCSNR J0512−6707 using available MIR imaging. The MIR maps used in this analysis are from 'Surveying the Agents of Galaxy Evolution' (SAGE) conducted by the *Spitzer* Space Telescope (Meixner et al. 2006). The 8 μm data is from the Infrared Array Camera (IRAC: Fazio et al. 2004) and the 24 μm band is from the Multiband Imaging Photometer for *Spitzer* (MIPS; Rieke et al. 2004). The pixel sizes are 0.6 arcsec for all IRAC wavelengths and 2.49 arcsec for 24 μm.

An image from the combined Hα (red), 8 μm (blue) and 24 μm (green) maps (Fig. 11) shows dense dust emission to the south which is clearly unrelated to the non-thermal SNR. Since the strong dust content at 8 μm dominates the other three shorter IRAC wavebands (see Reid 2014) they have not been shown in Fig. 11. The image shows a dust filament that appears to emanate from the strong emission to the south, proceed northwards, along the eastern border of the SNR, and then move across its northern boundary to the west. It therefore appears that the SNR is encased in dust, however, close examination from our line of sight shows a faint, diffuse cloud of 24 μm dust across the face of the SNR.

Although the 24 μm emission is spread in patches across the wider area, the contribution from 8 μm, shown in blue, is more pronounced in the densest areas of 24 μm dust. The 24 μm dust dominates the region immediately surrounding the SNR, however, patches of 8 μm emission are detectable to the east and north. Depending on our viewing angle, it is possible that the expanding SNR has impacted and cleared the 8 μm dust towards the SW and along the bright, outer shell at J2000 RA 05^{h}12^{m}30^s DEC −67°07 arcmin55 arcsec in the North/north-east. A dense region of 8 μm dust can be detected south of the southern perimeter of the

SNR which is also immediately south of what appears to be the radio centre (see Figs 8 and 9).

3.4 X-ray

The presence of X-ray emission is typical of most SNRs and is thought to arise from locations where shocks are converting explosive energy into cosmic ray energy. The radiative signatures of GeV and TeV particles indicate radio to X-ray synchrotron emission from relativisitc electrons as well as γ-ray emission from accelerated electrons and ions (Vink 2012). The X-ray source corresponding to MCSNR J0512−6707 may be found in Haberl & Pietsch (1999) *ROSAT* PSPC catalogue of X-ray sources in the LMC region, where it is listed as HP 483. The central coordinates of the source are J2000 RA 05^{h}12^{m}28.0^s DEC −67°07 arcmin27 arcsec, which is to within ∼4 arcsec of the centre of the remnant. The X-ray source extent from the central position is given as 8.5 ± 0.2 arcsec. The count rate (PSPC 0.1–2.4keV) is given as $3.33 \times 10^{-3} \pm 8.7 \times 10^{-4}$ and the hardness ratios are given as follows: H1 is 1.0 ± 0.38 and H2 is 0.04 ± 0.23. Further X-ray observations will be provided in Kavanagh et al. A&A, submitted.

4 CONCLUSION

We have discovered a new SNR in the LMC. This classification is based on combined evidence from optical imaging and spectroscopy, radio detections with a non-thermal spectral index and an X-ray counterpart. Both optical images shown in Fig. 1 clearly depict the shell structure and typical SNR-like morphology. The most distinguishing features are a series of bright loops across the northern boundary of the object. An [S II]/Hα ratio of 0.93 at one position across the outer shell suggests a typically strong shock. The presence and strength of the [O II] 3727, 3729 doublet and the [O III] 4959, 5007 lines are typical for SNRs. We note that the nitrogen [N II]/Hα ratios are low compared to those found in the spectra of Galactic SNRs, but this can be attributed to the fact that this element is less abundant in the LMC compared to the Galaxy (see Dopita 1979). Confusion with Wolf–Rayet ejecta can be ruled out as no WR star was detected in the vicinity and the MCELS survey showed Hα, not [O III], leading the emission. Assuming that the SNR is in the adiabatic stage, the Sedov–Taylor method was used to find the dynamical age of the remnant which may be between ∼2000 yr and ∼5000 yr. The expansion velocity will therefore be between ∼ 600 km s^{-1} and 1400 km s^{-1}.

There is a very strong radio source at the optical centre of the SNR. Radio-continuum data used in this project includes 36, 20, 6 and 3 cm mosaic images. These mosaic data were complemented by pointed observations of MCSNR J0512−6707 from our ATCA observations. The results of these measurements and those taken from the mosaics are listed in Table 4 and shown in Fig. 10. The resulting spectral index between 843 and 8640 MHz of $\alpha = -0.52 \pm 0.04$ is consistent with the typical SNR spectral index value of −0.5. Values derived for the magnetic field, from 124 to 184 μG, are relatively high for an SNR, and may suggest an age of only a few thousand years.

This SNR is surrounded by an area filled with dense and often compact dust, much of which may have been expelled outwards or cleared by the expanding shell of the SNR. Any contribution the SNR has made towards the local dust content is impossible to judge due to the large region of dense dust to the immediate south. An X-ray detection at the centre of the SNR was made through a previous survey (Haberl & Pietsch 1999).

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