

EVIDENCE FOR INFALLING GAS OF LOW ANGULAR MOMENTUM TOWARD THE L1551 NE KEPLERIAN CIRCUMBINARY DISK

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ABSTRACT

We report follow-up C¹⁸O(3–2) line observations of the Class I binary protostellar system L1551 NE with the Submillimeter Array in its compact and subcompact configurations. Our previous observations at a higher angular resolution in the extended configuration revealed a circumbinary disk exhibiting Keplerian motion. The combined data, with more extensive spatial coverage ($\sim 140\text{--}2000$ AU), verify the presence of a Keplerian circumbinary disk and reveal for the first time a distinct low-velocity ($\lesssim \pm 0.5$ km s⁻¹ from the systemic velocity) component that displays a velocity gradient along the minor axis of the circumbinary disk. Our simple model that reproduces the main features seen in the position–velocity diagrams comprises a circumbinary disk exhibiting Keplerian motion out to a radius of ~ 300 AU, beyond which the gas exhibits pure infall at a constant velocity of ~ 0.6 km s⁻¹. This velocity is significantly smaller than the expected free-fall velocity of ~ 2.2 km s⁻¹ onto the L1551 NE protostellar mass of $\sim 0.8 M_{\odot}$ at ~ 300 AU, suggesting that the infalling gas is decelerated as it moves into regions of high gas pressure in the circumbinary disk. The discontinuity in angular momenta between the outer infalling gas and the inner Keplerian circumbinary disk implies an abrupt transition in the effectiveness at which magnetic braking is able to transfer angular momentum outward, a result perhaps of the different plasma β values and the ionization fractions between the outer and inner regions of the circumbinary disk.

Key words: ISM: individual objects (L1551 NE) – ISM: molecules – stars: formation

Online-only material: color figures

1. INTRODUCTION

Compact and flattened circumstellar structures are commonly seen around T Tauri stars at the Class II stage of evolution for low-mass young stellar objects. Observations in molecular lines reveal that these structures exhibit Keplerian motion and therefore correspond to centrifugally supported circumstellar disks (Guilloteau & Dutrey 1998; Guilloteau et al. 1999; Simon et al. 2000; Piétu et al. 2003; Raman et al. 2006). Observations in dust continuum emission have unveiled the presence of compact ($\lesssim 600$ AU) and flattened structures also around protostars in both the Class 0 and more evolved Class I stages (Looney et al. 2000; Brinch et al. 2007; Enoch et al. 2009, 2011; Maury et al. 2010). In the case of Class 0 protostars, observations in molecular lines show primarily infalling motion only in their circumstellar disk-like structures (Takakuwa et al. 2007; Yen et al. 2010). When rotation also is detected, the rotational velocity is too small for the circumstellar disk-like structures to be centrifugally supported (Brinch et al. 2009; Yen et al. 2011, 2013). In contrast, an increasing number of studies have been successful at finding Keplerian motion in the circumstellar disks of Class I objects (Brinch et al. 2007; Lommen et al. 2008; Jørgensen et al. 2009; Tobin et al. 2012; Takakuwa et al. 2012; Yen et al. 2013; Hara et al. 2013). The masses of these objects’ circumstellar disks span the range $\sim 0.004\text{--}0.06 M_{\odot}$ and are therefore comparable to the masses of circumstellar disks around Class II objects. Their radii span the range $\sim 100\text{--}300$ AU, also comparable to the sizes of circumstellar disks around Class II objects. Recently, Tobin et al. (2012) reported the detection of a Keplerian disk around L1527 IRS, which these authors claim to be an object in transition from

Class 0 to I and therefore the youngest protostar known to exhibit a Keplerian circumstellar disk. The abovementioned results suggest that the transition from primarily infalling disk-like structures to Keplerian circumstellar disks around protostars occurs sometime during their evolution from Class 0 to Class I.

Although theoretical models have no problem forming disk-like structures around protostars, in many models these structures are not centrifugally supported and hence do not exhibit Keplerian motion (these disk-like structures are therefore commonly referred to as pseudodisks). Naively, if the infalling material from the surrounding dense core conserves angular momentum, eventually the rotational velocity of this material should match the Keplerian velocity and hence a centrifugally supported disk should form around the protostar. In many theoretical models, however, the inclusion of magnetic fields so effectively transports angular momentum outward through magnetic braking that no centrifugally supported structures ever form (e.g., Mellon & Li 2008, 2009; Li et al. 2011). From a series of three-dimensional MHD simulations of collapsing Bonnor–Ebert spheres, Machida & Matsumoto (2011) and Machida et al. (2011) were able to successfully produce centrifugally supported disks around protostars. The disks produced in their simulations, however, do not become centrifugally supported over radial sizes of $\gtrsim 100$ AU until nearly all of the material in their parental dense core has been exhausted. Furthermore, the disks have masses of a few to a few hundred times the mass of their central protostars and hence do not exhibit Keplerian motion. Both the mass and motion of the centrifugally supported disks produced in the simulations by Machida & Matsumoto (2011) and Machida et al. (2011) do not match those inferred for the Keplerian circumstellar disks of

Class I protostars. Nevertheless, the fact that Machida & Matsumoto (2011) and Machida et al. (2011) were able to successfully produce centrifugally supported disks around protostars may provide important insights in how Keplerian circumstellar disks form around protostars. Li et al. (2011) argue, however, that Machida & Matsumoto (2011) and Machida et al. (2011) do not properly handle the accumulation of magnetic flux at small radii because their simulations lack sufficient spatial resolution (i.e., numerical diffusivity), artificially weakening the efficiency of magnetic braking.

In a recent paper (Takakuwa et al. 2012, hereafter Paper I), we reported the detection of a Keplerian circumbinary disk around the Class I system L1551 NE. Located in the Taurus molecular cloud at a distance of ~ 140 pc (Elias 1978), L1551 NE exhibits two 3.6-cm radio continuum sources (comprising free-free emission from protostellar jets) with a projected separation of ~ 70 AU at a position angle (P.A.) of $\sim 300^\circ$ (Reipurth et al. 2002). Reipurth et al. (2002) referred to the south-eastern source as “Source A” and the northwestern source as “Source B,” a nomenclature that we have retained. Near-infrared (NIR) observations revealed that Source A drives a pair of collimated [Fe II] jets along the northeast to southwest direction at a P.A. of $\sim 60^\circ$ (Reipurth et al. 2000, 2002; Hayashi & Pyo 2009). Source B is located at the origin of an extended (~ 2000 AU) NIR reflection nebula corresponding to an outflow cavity (Reipurth et al. 2000, 2002; Hayashi & Pyo 2009). Source B must therefore have driven a stronger outflow in the recent past that carved out the observed outflow cavity. Our observations of L1551 NE reported in Paper I were made with the Submillimeter Array (SMA)⁵ in its extended configuration in the 0.9 mm continuum as well as the ^{13}CO (3–2) and C^{18}O (3–2) lines. These observations revealed compact (unresolved) dust components associated with each binary component suggestive of circumstellar disks, as well as a circumbinary disk of radius ~ 300 AU exhibiting Keplerian motion. The mass (in both molecular gas and dust) of the circumbinary disk, estimated from the 0.9 mm dust continuum, is $\sim 0.03 M_\odot$. From model fitting to the C^{18}O velocity channel maps, we inferred a disk inclination of $\sim 62^\circ$, P.A. for its major axis of $\sim 167^\circ$, and total mass for the enclosed binary protostellar components of $\sim 0.8 M_\odot$; i.e., the circumbinary disk is about an order of magnitude less massive than the protostellar system contained within. Evidently, by or during the Class I stage, Keplerian disks have formed or can form not only around individual protostars, but also around the binary components of a protostellar system.

Due both to its relative proximity and the relatively large size of its circumbinary disk, L1551 NE presents a good opportunity to study the kinematics of material just beyond the centrifugally supported disk of a protostellar system. In this way, we hope to improve models of how such disks form. In this paper, we report follow-up observations of L1551 NE with the SMA in the subcompact and compact configurations that cover spatial scales as large as ~ 2000 AU, which in our previous observations (Paper I) were strongly if not completely resolved out by the interferometer. Our goals are to better define the outermost extent of the circumbinary disk exhibiting Keplerian motion (hereafter, the Keplerian radius) and to infer the dynamics of material immediately beyond this portion of the disk. In Section 2, we describe our observations and data reduction. In Section 3, we present our results for the spatial and velocity

distribution of the C^{18}O (3–2) emission from the new data, as well as the combined new and old (from Paper I) data. In Section 4, we model the observed gas kinematics to infer the centrifugally supported extent of the circumbinary disk. We also report the detection of infalling material with little, if any, angular momentum beyond the Keplerian radius and discuss the implications of our findings for the formation of the Keplerian circumbinary disk in L1551 NE. In Section 5, we provide a concise summary of the results and our interpretations and offer a few thoughts for future work.

2. OBSERVATIONS AND DATA REDUCTION

We observed L1551 NE in the C^{18}O $J = 3-2$ line (rest frequency of 329.3305453 GHz) with the SMA (Ho et al. 2004) in its compact configuration on 2011 December 28 and its subcompact configuration on 2012 January 14. In Table 1, we summarize the relevant parameters of the telescope and the calibrators used in our observations. The frequency resolution employed in these observations, corresponding to a velocity resolution of ~ 0.185 km s $^{-1}$, is a factor of two higher than that used in our previous observations when the SMA was in its extended configuration (Paper I). The raw visibility data were calibrated and flagged with MIR, which is an IDL-based data reduction package (Scoville et al. 1993). We then used the software package Miriad (Sault et al. 1995) to Fourier transform and CLEAN the calibrated visibility data to produce the final images. The images made from the newly acquired data in the compact and subcompact configurations added together are henceforth referred to as the “new” images. We also combined our newly acquired data with our previous data taken in the extended configuration (Paper I) to produce “combined” images. The resultant image parameters are listed in Table 2. The minimum projected baseline length is ~ 6.2 k λ , so that 10%–50% of the peak flux can be recovered for a Gaussian emission distribution with an FWHM of $15''$ – $27''$ (Wilner & Welch 1994), and a larger fraction still for a Gaussian emission distribution with a smaller FWHM.

3. RESULTS

Figure 1 shows the combined C^{18}O (3–2) integrated intensity map of L1551 NE. As in Paper I, we find an elongated feature along the northwest to southeast direction centered approximately at the position of Source A. The major and minor axes of this elongated feature match well with the corresponding axes of the circumbinary disk reported in Paper I. As explained in that paper, this circumbinary disk has its major axis at a P.A. of 347° and minor axis at 77° , as indicated by the dashed lines in Figure 1. The major axis of the circumbinary disk is approximately perpendicular to the axis of the [Fe II] jets driven by Source A (Hayashi & Pyo 2009). In comparison with the circumbinary disk reported in Paper I, the elongated feature detected here has an outer dimension of ~ 1000 AU \times 800 AU that is significantly larger than the dimensions of the circumbinary disk of ~ 600 AU \times 300 AU (measured from data taken only with the extended configuration).

Figure 2 shows the combined velocity channel maps in C^{18}O (3–2). From the symmetry in velocity space evident in these maps and the position–velocity (P – V) diagrams described below, we hereafter adopt a systemic velocity for L1551 NE of $V_{\text{LSR}} = 6.9$ km s $^{-1}$. Over the blueshifted range ($V_{\text{LSR}} = 4.2$ – 6.1 km s $^{-1}$), the C^{18}O emission is located predominantly to the north (hereafter referred to as the

⁵ The SMA is a joint project between the Smithsonian Astrophysical Observatory and the Academia Sinica Institute of Astronomy and Astrophysics and is funded by the Smithsonian Institution and the Academia Sinica.

Table 1
Parameters for the SMA Observations of L1551 NE

Parameter	Value	
	2011 December 28	2012 January 14
Number of antennae	8	7
Configuration	Compact	Subcompact
Flux calibrator	Uranus	Callisto
Gain calibrator	0423–013, 0530+135	
Flux (0423–013 upper)	2.09 Jy	1.99 Jy
Flux (0423–013 lower)	2.08 Jy	1.99 Jy
Flux (0530+135 upper)	0.96 Jy	1.02 Jy
Flux (0530+135 lower)	0.95 Jy	1.08 Jy
Passband calibrator	3c84, 3c279, Uranus	bllac, 3c279, Callisto, Mars
System temperature double sideband	~250–1000 K	
Absolute flux uncertainty	~30%	

Table 2
Parameters of the C¹⁸O (3–2) Image Cube Taken with the SMA Subcompact, Compact, and Extended Configurations

Parameter	All the Config.	Sub. + Compact Config.
($\alpha_{J2000.0}$, $\delta_{J2000.0}$)	(04 ^h 31 ^m 44 ^s .47, 18°08′32″.2)	
Primary beam HPBW	~37″	
Synthesized beam HPBW	1″.36 × 1″.06 (P.A. = –86°)	2″.50 × 2″.20 (P.A. = –61°)
Frequency resolution	406.25 (kHz) ~0.37 (km s ^{–1})	203.125 (kHz) ~0.185 (km s ^{–1})
Baseline coverage	6.2–249.0 (k λ)	6.2–84.5 (k λ)
Conversion factor	7.82 (K/J beam ^{–1})	2.05 (K/J beam ^{–1})
rms noise level	0.078 (Jy beam ^{–1})	0.18 (Jy beam ^{–1})

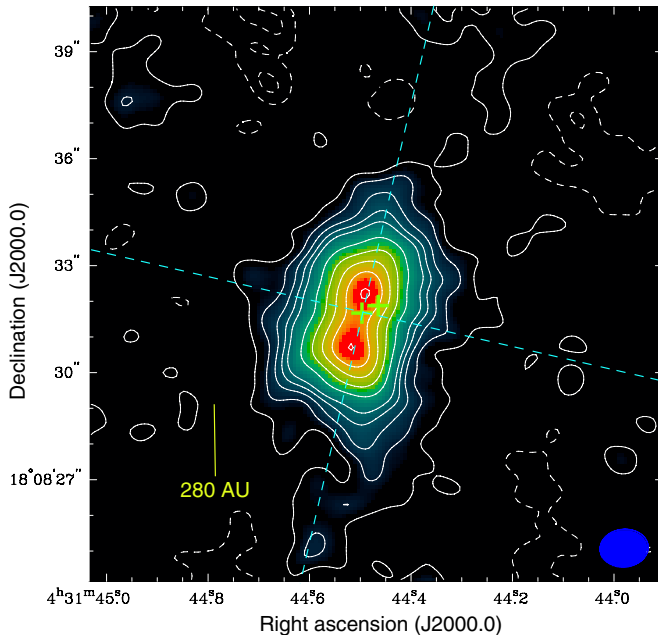


Figure 1. Integrated C¹⁸O (3–2) intensity (moment 0) map made by combining the newly acquired data in the subcompact and compact configurations with the previously published data (Paper I) in the extended configuration of the SMA. Contour levels are from 2σ in steps of 2σ to 10σ , and then in steps of 5σ ($1\sigma = 0.119 \text{ Jy beam}^{-1} \text{ km s}^{-1}$). The velocities used to make this map span $V_{\text{LSR}} = 3.85\text{--}10.14 \text{ km s}^{-1}$. Crosses indicate the positions of the binary components of L1551 NE (from Reipurth et al. 2002), with Source A located to the southeast and Source B located to the northwest. Tilted dashed lines denote the major (P.A. = 347°) and minor (P.A. = 77°) axes of the Keplerian circumbinary disk discovered in Paper I, which pass through Source A. The filled ellipse at the bottom-right corner corresponds to the synthesized beam as plotted at FWHM ($1''.36 \times 1''.06$ at P.A. = -86°).

(A color version of this figure is available in the online journal.)

“high-velocity blueshifted” component), whereas over the redshifted range ($7.9\text{--}9.4 \text{ km s}^{-1}$), the emission is located to the south (the “high-velocity redshifted” component) of the protostellar system. Over the redshifted range of $7.2\text{--}7.5 \text{ km s}^{-1}$ (the “low-velocity redshifted” component), there is strong emission to the south, as is seen for the high-velocity redshifted component, but also relatively weak and extended emission to the north. Figure 3 shows velocity channel maps over a smaller velocity range of $6.6\text{--}7.2 \text{ km s}^{-1}$ but at a higher velocity resolution (0.185 km s^{-1}) made from just the new data. In these higher velocity resolution channel maps, the redshifted emission close to the systemic velocity can be seen to be more strongly weighted toward the east of the protostellar system whereas the blueshifted emission is more strongly weighted toward the west. Thus, there appears an east (red) to west (blue) velocity gradient at low velocities close to the systemic velocity.

Figure 4 shows the combined (left column) and new maps (right column) integrated over the high-velocity blueshifted and redshifted components (top row), the low-velocity blueshifted component (middle row), and the low-velocity redshifted component (bottom row). In the high-velocity components, the blueshifted and redshifted emission are well separated and located to the north and south of the protostellar system, respectively. The orientation of these blueshifted and redshifted components and hence the orientation of their velocity gradient match well with the orientation of the major axis of the circumbinary disk, suggesting that the high-velocity components primarily trace the Keplerian circumbinary disk as found in Paper I. In the low-velocity components, blueshifted emission is seen toward the north of the binary components as in the high-velocity components; on the other hand, redshifted emission is seen not just to the south but also to the north of the binary components. Along the minor axis, the intensity of the low-velocity redshifted emission is weighted toward the east of the protostellar system and that of the low-velocity blueshifted

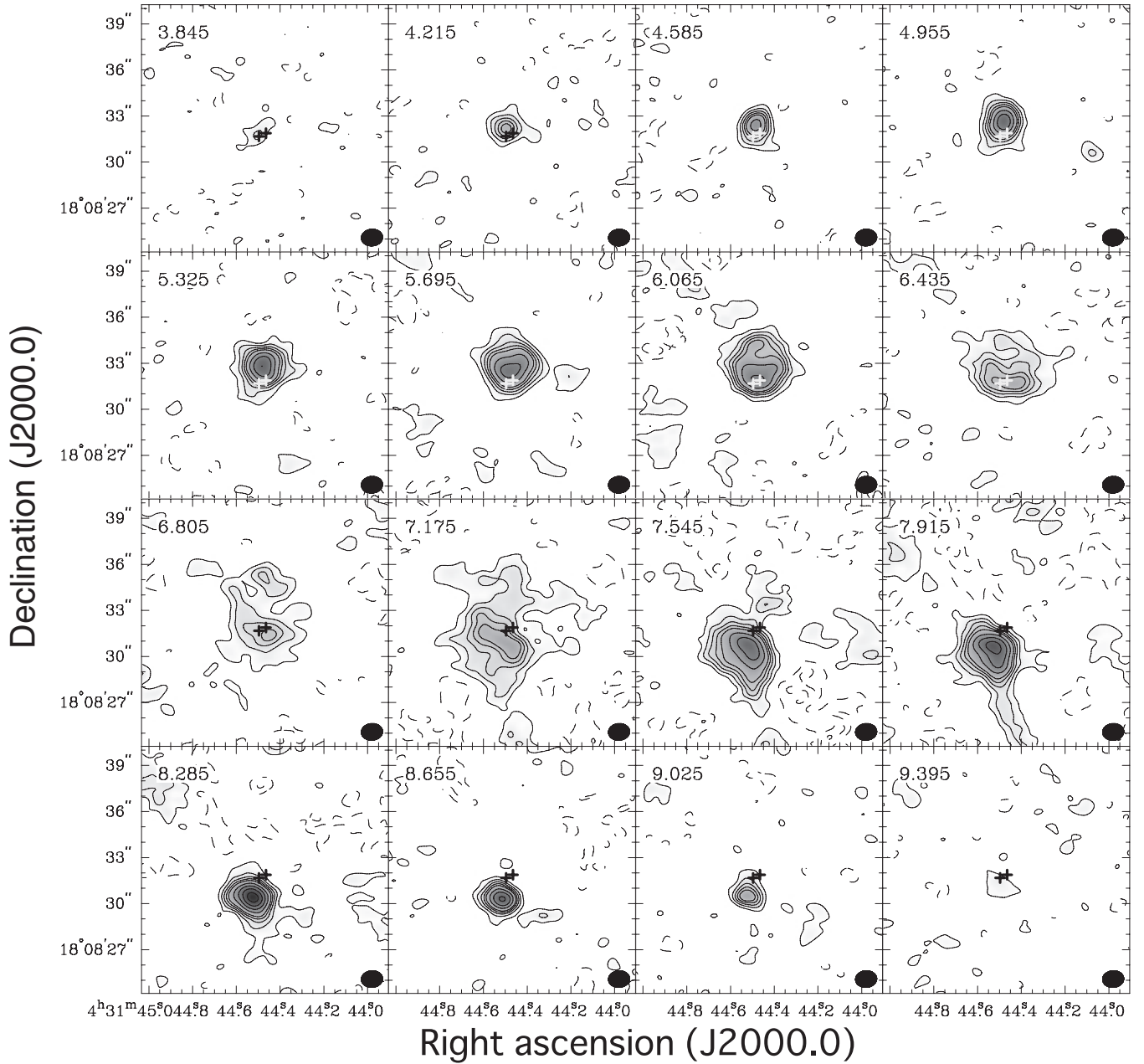


Figure 2. Velocity channel C^{18}O (3–2) maps made with the newly acquired and previously published (Paper I) data combined. Contour levels are from 2σ in steps of 2σ to 10σ , and then in steps of 4σ ($1\sigma = 0.081 \text{ Jy beam}^{-1}$). The local standard of rest (LSR) velocity is indicated at the top-left corner of each panel. Crosses indicate the positions of the binary components of L1551 NE, as in Figure 1. The filled ellipse at the bottom-right corner of each panel corresponds to the synthesized beam as plotted at $\text{FWHM} (1''.36 \times 1''.06 \text{ at P.A.} = -86^\circ)$.

emission is weighted toward the west, indicating an east (red) to west (blue) velocity gradient along the minor axis at velocities close to the systemic velocity, as discussed earlier. This velocity gradient along the minor axis is clearly seen in the $P-V$ diagram of C^{18}O (3–2), as we describe next.

Figure 5 (left column) shows the $P-V$ diagrams in C^{18}O , derived from the combined images, along the major (upper row) and minor (lower row) axes of the circumbinary disk and passing through the position of Source A. Along the major axis, the high-velocity blueshifted and redshifted components are individually separated and located to the north and south, respectively, of Source A. The emission at higher velocities is concentrated closer to Source A than that at lower velocities, a characteristic feature of Keplerian disks illustrated by the blue

curves corresponding to the Keplerian rotation curve derived in Paper I for the circumbinary disk. In addition to the features tracing Keplerian rotation, there is also a spatially extended redshifted feature around the systemic velocity toward the north of the Source A (not seen previously in Paper I); this feature arises from the low-velocity redshifted component toward the north of the binary components mentioned above. Along the minor axis, the $P-V$ diagram exhibits a clear velocity gradient at velocities relatively close to the systemic velocity (again not previously seen in Paper I), as indicated by the arrows in the lower left panel of Figure 5. In contrast, no velocity gradient is evident along the minor axis at higher velocities. The velocity gradient seen at relatively low velocities along the minor axis therefore reflects the motion of the eastern low-velocity

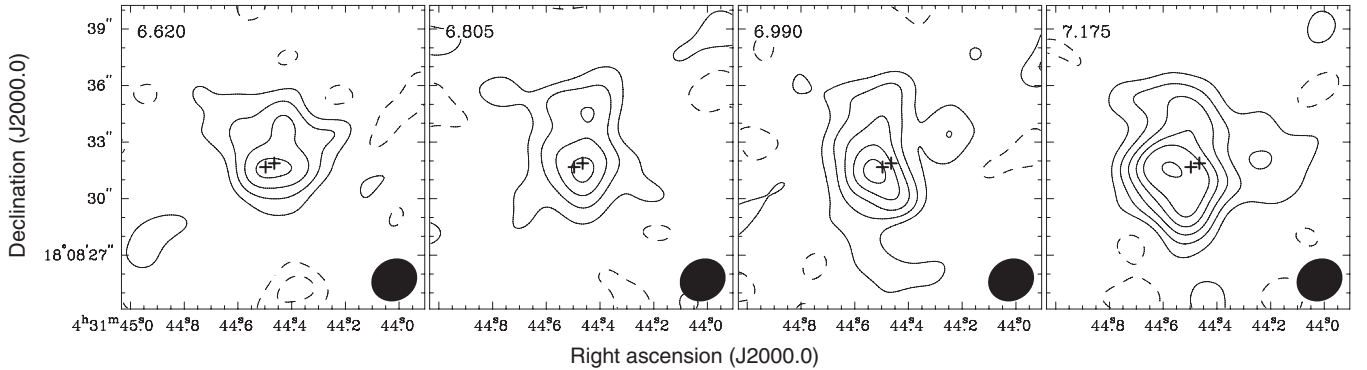


Figure 3. Velocity channel C^{18}O (3–2) maps made with the newly acquired data only, which have a factor of two higher velocity resolution than the previously published data (Paper I), at velocities close to the systemic velocity (6.9 km s^{-1}). Contour levels are from 2σ in steps of 2σ to 10σ , and then 14σ ($1\sigma = 0.18 \text{ Jy beam}^{-1}$). The LSR velocity is indicated at the top-left corner of each panel. Crosses indicate the positions of the binary components of L1551 NE, as in Figure 1. The filled ellipse at the bottom-right corner of each panel corresponds to the synthesized beam as plotted at FWHM ($2''.50 \times 2''.20$ at P.A. = -61°).

redshifted and western low-velocity blueshifted components. For purely Keplerian rotation, no velocity gradient should be evident along the minor axis, and indeed no velocity gradient is seen along the minor axis for the high-velocity components. The observed velocity gradient along the minor axis for emission at relatively low velocities, seen here for the first time from observations sensitive to larger-scale structures than in Paper I, therefore indicates motion not related to Keplerian motion. In the next section, we will discuss the nature of this newly discovered component.

4. DISCUSSION

4.1. Nature of the Low-velocity Components

As described in the previous section, we find that the C^{18}O (3–2) emission surrounding L1551 NE exhibits two distinct velocity components: (1) high-velocity blueshifted and redshifted components that are more spatially extended along the north–south than east–west directions, and which show a velocity gradient along the more spatially extended north–south direction. These components trace a Keplerian circumbinary disk as identified in Paper I. (2) low-velocity blueshifted and redshifted components that also are more spatially extended along the north–south than the east–west directions, but which, unlike the high-velocity blueshifted and redshifted components, show a velocity gradient along the less spatially extended east–west direction. Below, we consider two alternative explanations and offer our opinion for the likely nature of the low-velocity components.

One possible interpretation for the low-velocity components is a molecular outflow. From observations with the James Clerk Maxwell telescope (JCMT), Moriarty-Schieven et al. (1995) found that the CO (3–2) emission around L1551 NE exhibits blueshifted and redshifted line wings (at velocities $\gtrsim 2.5 \text{ km s}^{-1}$ from the systemic velocity) in spectra taken toward the west and east of the protostellar system, respectively. These authors attributed the emission in the line wings to a molecular outflow driven by L1551 NE. Furthermore, near-IR [Fe II] observations of L1551 NE show knotty jets driven by Source A, which are blueshifted to the southwest and redshifted to the northeast (Reipurth et al. 2000, 2002; Hayashi & Pyo 2009). Thus, the line wings in CO (3–2), as well as the jets driven by Source A, share the same sense in velocity gradient in the sky as the low-velocity C^{18}O (3–2) emission reported above. If the low-velocity blueshifted and redshifted C^{18}O (3–2) components indeed trace

a molecular outflow, then based on their projected velocity ($\sim 0.5 \text{ km s}^{-1}$; see black arrows in Figure 5) and projected separation from Source A ($\sim 2''$, corresponding to $\sim 280 \text{ AU}$) and assuming that the outflow emerges perpendicular to the Keplerian circumbinary disk (inclination of $i \sim 62^\circ$; see Paper I), the deprojected outflow velocity is then $v_{\text{flow}} \sim 0.5 \text{ km s}^{-1} / \cos i \sim 1.0 \text{ km s}^{-1}$ and the deprojected spatial separation of the individual outflow components $l_{\text{flow}} \sim 280 \text{ AU} / \sin i \sim 320 \text{ AU}$ from Source A. By comparison, at their separation from the protostellar system, the escape velocity is

$$v_g = \sqrt{\frac{2GM_\star}{l_{\text{flow}}}}, \quad (1)$$

where G is the gravitational constant and M_\star is the mass of the protostellar system ($= 0.8 M_\odot$; Paper I). We find $v_g \sim 2.1 \text{ km s}^{-1}$, a factor two higher than the deprojected outflow velocity mentioned above of $v_{\text{flow}} \sim 1.0 \text{ km s}^{-1}$. The low-velocity components are therefore bound to the protostellar system. Indeed, the low-velocity components reach much lower blueshifted and redshifted velocities ($\lesssim \pm 0.5 \text{ km s}^{-1}$ from the systemic velocity) than the CO (3–2) outflow observed with the JCMT, even though the JCMT observations probe regions much further away from L1551 NE than those spanned by the low-velocity components. Furthermore, contrary to expectations for a molecular outflow driven along the east–west directions, the low-velocity components are even more spatially extended along the north–south than the east–west directions.

Alternatively, we consider the possibility that the low-velocity components correspond to gas coplanar with but lying just beyond the Keplerian radius of the circumbinary disk. As described in Section 3, the major axis of the circumbinary disk is orthogonal to the projected axes of the [Fe II] jets from Source A; if indeed the plane of the circumbinary disk is perpendicular to the axis of the jets, the western part of the disk plane would then correspond to the far side and the eastern part of the disk plane to the near side. Correspondingly, the western low-velocity blueshifted component would then be located on the far side and the eastern low-velocity redshifted component would be located on the near side. The low-velocity components would therefore be infalling toward the circumbinary disk along the disk plane. The infall velocity, v_{inf} , can be roughly estimated from the velocity gradient seen for the low-velocity components along the minor axis of the circumbinary disk in the P – V diagram of Figure 5. At the opposing locations of the

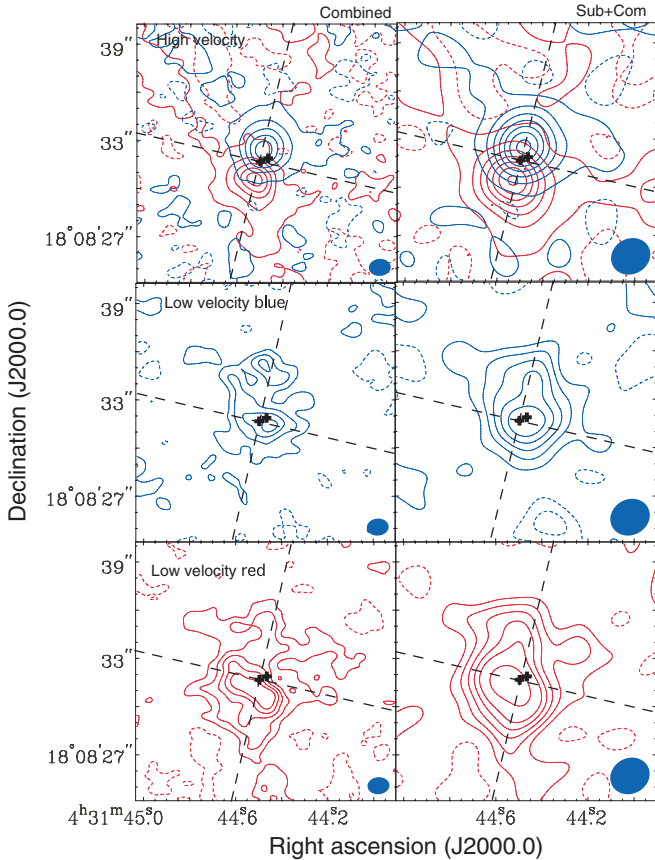


Figure 4. Integrated C^{18}O (3–2) intensity maps over selected velocity ranges corresponding to the high-velocity (upper row), low-velocity blueshifted (middle row), and low-velocity redshifted (lower row) components (see the text). The left column shows images made with the newly acquired and previous published data (Paper I) combined, and the right column just the newly acquired data that provide higher sensitivity to large-scale structures. Blueshifted emission is plotted in blue contours and redshifted emission in red contours. In the left panels, the LSR velocity range used to make the high-velocity blueshifted component spans $3.66\text{--}6.62\text{ km s}^{-1}$, the high-velocity redshifted component spans $7.36\text{--}10.32\text{ km s}^{-1}$, the low-velocity blueshifted component spans $6.62\text{--}6.99\text{ km s}^{-1}$, and the low-velocity redshifted component spans $6.99\text{--}7.36\text{ km s}^{-1}$. In the right panels, the velocity range used to make the high-velocity blueshifted component spans $3.7525\text{--}6.5275\text{ km s}^{-1}$, the high-velocity redshifted component spans $7.2675\text{--}10.0425\text{ km s}^{-1}$, the low-velocity blueshifted component spans $6.5275\text{--}6.8975\text{ km s}^{-1}$, and the low-velocity redshifted component spans $6.8975\text{--}7.2675\text{ km s}^{-1}$. Contour levels are plotted from 2σ in steps of 8σ for the high-velocity component ($1\sigma = 0.029\text{ Jy beam}^{-1}$ and $0.0465\text{ Jy beam}^{-1}$, respectively, in the left and right panels) and 2σ in steps of 2σ to 10σ and then 14σ for the low-velocity component ($1\sigma = 0.081\text{ Jy beam}^{-1}$ and $0.127\text{ Jy beam}^{-1}$, respectively, in the left and right panels). Crosses and dashed lines are the same as in Figure 1. The filled ellipse at the bottom-right corner of each panel corresponds to the synthesized beam as plotted at the FWHM, with sizes as specified in Figure 2 for the left column and in Figure 3 for the right column.

(A color version of this figure is available in the online journal.)

low-velocity components, the velocity shift across the minor axis is $\sim 1.0\text{ km s}^{-1}$ (see the arrows in Figure 5). At an inclination of $i \sim 62^\circ$ for the circumbinary disk, the infall velocity along the disk plane is therefore $v_{\text{inf}} \sim (1.0\text{ km s}^{-1}/2)/\sin 62^\circ \sim 0.6\text{ km s}^{-1}$.

To see if the above interpretation offers a good fit to the observed P – V diagrams, we have generated P – V diagrams based on a simple toy model to compare with the observed P – V diagrams. Our model comprises a geometrically thin disk, as was adopted to model the Keplerian circumbinary disk described in

Paper I, but with a velocity field now described by

$$v_{\text{rot}}(r) = \sqrt{\frac{GM_\star}{r}} \quad \text{and} \quad v_{\text{rad}}(r) = 0 \quad \text{for} \quad r \leq r_{\text{kep}}, \quad (2a)$$

$$v_{\text{rot}}(r) = 0 \quad \text{and} \quad v_{\text{rad}}(r) = v_{\text{inf}} \quad \text{for} \quad r > r_{\text{kep}}, \quad (2b)$$

where r is the radius along the disk, $v_{\text{rot}}(r)$ is the rotational velocity of the disk (i.e., the portion corresponding to the Keplerian circumbinary disk), $v_{\text{rad}}(r)$ is the radial (infall) velocity of the disk (i.e., the portion corresponding to the low-velocity components), r_{kep} is the Keplerian radius, and v_{inf} is an adopted constant infall velocity beyond the Keplerian radius. As in Paper I, we adopt $M_\star = 0.8 M_\odot$, $i = 62^\circ$, a disk P.A. of $\theta = 167^\circ$ with the dynamical center at Source A, and an internal velocity dispersion $\sigma_{\text{gas}} = 0.4\text{ km s}^{-1}$. To model the intensity distribution of the disk, we adopt an intensity profile that resembles the intensity distribution of the C^{18}O emission as derived from the moment 0 map (the C^{18}O emission shows no evidence for candidate circumstellar disks associated with the two binary components, as seen in the continuum at 0.9 mm and reported in Paper I; we therefore believe that the observed C^{18}O emission is dominated by the circumbinary disk component). That is, we fit two-dimensional Gaussian structures to the observed moment 0 map of Figure 1 to represent the intensity distribution of the disk; we found that a minimum of three Gaussian components were required to produce a satisfactory fit, comprising two relatively compact Gaussians each centered on the two peaks seen in the moment 0 map, along with a much broader Gaussian to capture more extended emission surrounding the two peaks. By varying the Keplerian radius (r_{kep}) and the infall velocity beyond the Keplerian radius (v_{inf}), we found from trial and error a model P – V diagram that we judged (by eye) to provide a satisfactory fit to the observed P – V diagrams. To generate the model P – V diagrams, first each trial velocity field for the disk was converted into the line-of-sight velocity on the disk plane and then model velocity channel maps were created from the line-of-sight velocity map, the model moment 0 map, and the assumed internal velocity dispersion (σ_{gas}). The model P – V diagrams were then generated from the model velocity channel maps.

The model described above ignores a potential envelope component (corresponding to the parental dense core of L1551 NE) in which the disk component is embedded. If this envelope component is not completely recovered in our SMA observations, it introduces artifacts in the maps constructed from the available data and the velocity field derived from these maps (i.e., the P – V diagrams). To simulate the effect such an envelope would have on our model P – V diagrams, we added an envelope component to the disk component described above so that our final model comprises a disk embedded in a spherically symmetric envelope. Using the uv coverage attained in our SMA observations, we made mock observations of the model velocity channel maps having different r_{kep} and v_{inf} for the disk component as described above and the parameters for the envelope component as described below to generate model visibilities. We then Fourier transformed and CLEANED the model visibilities to produce simulated model velocity channel maps. These maps are then used to derive simulated model P – V diagrams. We decided on the parameters of the model envelope component in the following manner. A C^{18}O (3–2) spectrum

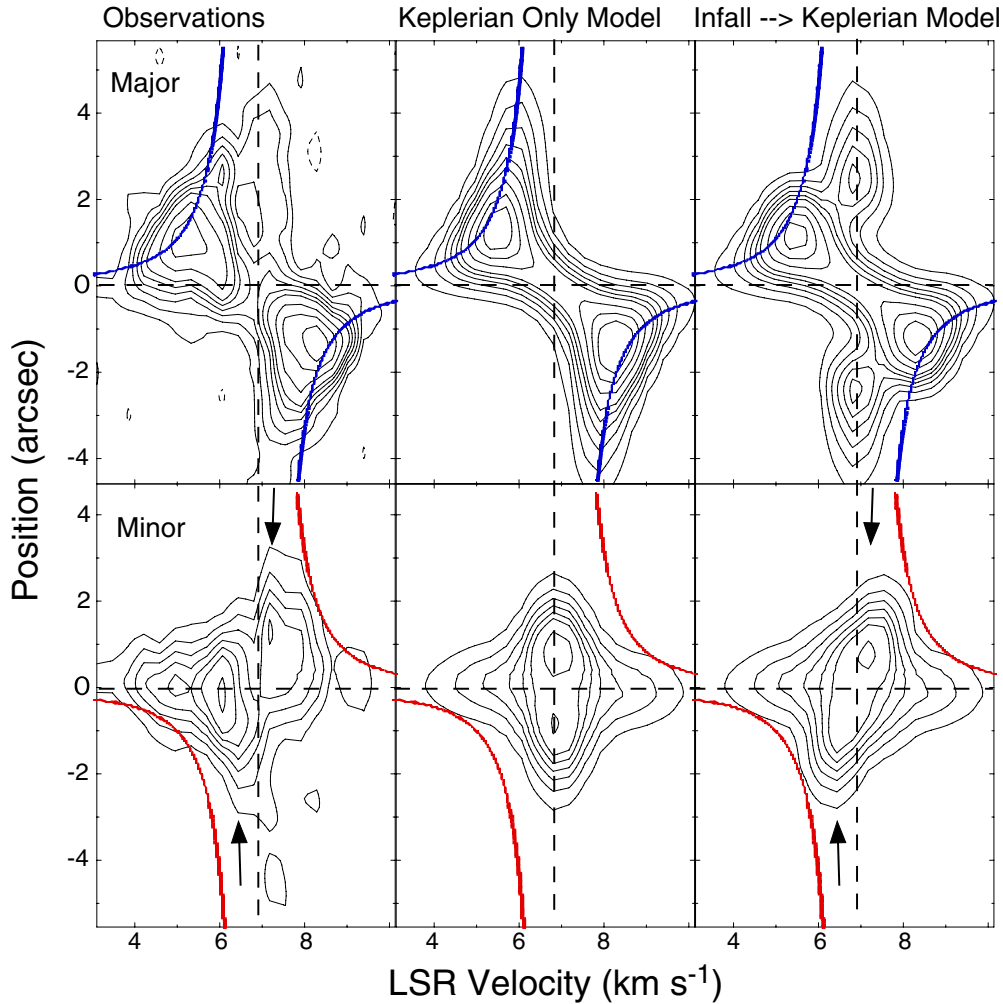


Figure 5. $P-V$ diagrams along P.A. = 347° in the upper row and P.A. = 77° in the lower row, centered on Source A. These position angles correspond to the major and minor axes, respectively, of the Keplerian circumbinary disk reported in Paper I. $P-V$ diagrams in the left column are derived from the C^{18}O (3–2) maps made from the newly acquired data combined with the previously published (Paper I) data. $P-V$ diagrams in the middle panels are derived from our model maps comprising a Keplerian circumbinary disk around a central stellar mass of $0.8 M_\odot$ and observed at an inclination angle of 62° . $P-V$ diagrams in the right panel are from our model maps comprising a circumbinary disk exhibiting Keplerian motion out to a radius of 300 AU and infall beyond this radius at a constant velocity of $\sim 0.6 \text{ km s}^{-1}$. Both models include an extended envelope component with properties as described in the text, and the resultant model images sampled with the uv coverage attained in our combined observations with the SMA. Contour levels are from 2σ in steps of 2σ to 12σ , and then in steps of 4σ ($1\sigma = 0.081 \text{ Jy beam}^{-1}$). Vertical dashed lines indicate the systemic velocity for L1551 NE of 6.9 km s^{-1} . Blue curves show the Keplerian rotation curve for a central stellar mass of $0.8 M_\odot$ and observed at an inclination angle of 62° , and red curves show free-fall in the plane of the Keplerian circumbinary disk onto a central stellar mass of $0.8 M_\odot$. Black arrows denote the velocity gradient detected along the minor axis.

(A color version of this figure is available in the online journal.)

taken with the Caltech Submillimeter Observatory (CSO) toward the center of L1551 NE exhibits two distinct components: one centered at $V_{\text{LSR}} \sim 6.7 \text{ km s}^{-1}$ that has a narrower velocity width of $\sim 0.68 \text{ km s}^{-1}$ and the other centered at $V_{\text{LSR}} \sim 7.0 \text{ km s}^{-1}$ that has a broader velocity width of $\sim 2.2 \text{ km s}^{-1}$ (Fuller & Ladd 2002). The central velocity and velocity range spanned by the broader component are similar to those of the C^{18}O (3–2) emission that we mapped with the SMA. Indeed, a comparison of the SMA C^{18}O (3–2) spectrum with the CSO C^{18}O (3–2) spectrum shows that virtually all of the flux in the CSO spectrum is recovered in our SMA observations in the LSR velocity range $\sim 3.9\text{--}5.4 \text{ km s}^{-1}$ and $\sim 8.7\text{--}9.4 \text{ km s}^{-1}$ (although the poor signal-to-noise ratio of the CSO spectrum in the line wings prevents us from precisely estimating the amount of recovered flux). In contrast, around the systemic velocity, we miss a significant amount of flux present in the CSO spectrum in our SMA observations. In the lower blueshifted velocity range

$6.4\text{--}6.8 \text{ km s}^{-1}$, the amount of the recovered flux is only 7%–8% of the CSO flux, whereas in the lower redshifted velocity range $7.2\text{--}7.5 \text{ km s}^{-1}$, the amount of recovered flux is 30%–40%. Much the missing flux corresponds to the narrower component in the CSO spectrum described above; this component must therefore correspond to a spatially extended envelope. Unfortunately, there are no available single-dish maps of the C^{18}O (3–2) emission around L1551 NE and so the structure and kinematics of this envelope are unclear. Observations of L1551 NE in the $850 \mu\text{m}$ dust continuum emission with SCUBA on the JCMT show that the size of the surrounding dusty envelope is $\sim 9''.8 \times 8''.8$ (Moriarty-Schieven et al. 2006). We therefore anticipate that the extent of the envelope in C^{18}O (3–2) is unlikely to be much larger than the CSO beam of $\sim 23''$. For the envelope component in our model, we have therefore adopted a spherically symmetric Gaussian profile with a cross-sectional size at FWHM that is the same as that of the CSO beam, together with a central

velocity ($V_{\text{LSR}} = 6.7 \text{ km s}^{-1}$), line width (0.68 km s^{-1}), and total intensity (55.5 Jy) having the same values as the corresponding parameters for the narrower component in the CSO C^{18}O spectrum.

Figure 5 (middle column) shows the simulated model $P-V$ diagrams along the major (upper row) and minor (lower row) axes with $r_{\text{kep}} = \infty$ so that the model disk exhibits Keplerian motion only (labeled the “Keplerian only model”). In the right column, the corresponding simulated model $P-V$ diagrams are shown, but now with $r_{\text{kep}} = 300 \text{ AU}$ and $v_{\text{inf}} = 0.6 \text{ km s}^{-1}$ so that the model disk exhibits infall (with no rotation) beyond the Keplerian radius (labeled the “infall \rightarrow Keplerian model”). (Recall that the intensity profile of the observed moment 0 map constrains the outer radius of the model disk that contributes significantly to the emission.) We also made simulated model velocity channel maps without the envelope component to check whether the inclusion of this component changes the simulated model $P-V$ diagrams compared with those shown in Figure 5, but did not find any important differences. For comparison, the best-fit Keplerian rotation curve derived in Paper I is drawn in all the $P-V$ diagrams along the major axis (blue curves in Figure 5) and the curve for free-fall onto a central system mass of $M_{\star} = 0.8 M_{\odot}$ is drawn in all the $P-V$ diagrams along the minor axis (red curves). In the $P-V$ diagram along the major axis, the model that invokes only Keplerian rotation in the disk (Keplerian only model) shows that the blueshifted and redshifted components lie at distinct opposite quadrants in the $P-V$ plane, such that emission at higher velocities is located closer to the center than emission at lower velocities; these characteristics are similar to those seen in the $P-V$ diagrams for the observed high-velocity blueshifted and redshifted components. This model, however, cannot reproduce the observed spatially extended low-velocity component to the north of Source A in the $P-V$ diagram along the major axis. Furthermore, in the $P-V$ diagram along the minor axis, the model incorporating only a Keplerian disk does not reproduce the observed velocity gradient at velocities relatively close to the systemic velocity. If we truncate the Keplerian disk at $r_{\text{kep}} \sim 300 \text{ AU}$ and add an outer infalling component along the disk plane with $v_{\text{inf}} \sim 0.6 \text{ km s}^{-1}$ (infall \rightarrow Keplerian model), we can reproduce the extended low-velocity component in the $P-V$ diagram along the major axis as well as the velocity gradient in the $P-V$ diagram along the minor axis at velocities relatively close to the systemic velocity.

In summary, our simple toy model of a circumbinary disk exhibiting a combination of Keplerian motion in its inner region and infall beyond its Keplerian radius provides an excellent fit to the gross spatial and kinematic morphology of the observed C^{18}O emission as captured in the $P-V$ diagrams (irrespective of whether or not the disk is embedded in a spatially extended envelope). We therefore believe that the low-velocity components newly detected here in observations sensitive to larger scale structures than those in Paper I most likely trace infalling gas located just beyond the Keplerian circumbinary disk first detected in Paper I. We note that the model (constant) infall velocity beyond the Keplerian radius of $v_{\text{inf}} \sim 0.6 \text{ km s}^{-1}$ (arrows in Figure 5) is smaller than the expected free-fall velocity (red curves in Figure 5) onto a central stellar mass (for both binary components combined) of $M_{\star} = 0.8 M_{\odot}$ over the radial range from the protostellar system (Source A) plotted in Figure 5; i.e., at a radius of $\sim 300 \text{ AU}$ (the Keplerian radius), the predicted free-fall velocity is $\sim 2.2 \text{ km s}^{-1}$. Instead, the model infall velocity of $\sim 0.6 \text{ km s}^{-1}$ at $r = 300 \text{ AU}$ implies a central stellar mass of only $\sim 0.06 M_{\odot}$, which is an order of magnitude

smaller than the formal value for the mass of the protostellar system estimated from the Keplerian motion of the circumbinary disk of $0.8_{-0.4}^{+0.6} M_{\odot}$ (uncertainty corresponding to the FWHM of the χ^2 curve; see Paper I). The smaller-than-expected infall velocity may indicate that the infalling material is decelerated (presumably by gas pressure, if not also by turbulence) as it approaches the Keplerian disk. On the other hand, we emphasize that, in our toy model, we did not include any rotation in the infalling region. We have attempted a number of models that connect the rotational profile of the Keplerian disk at r_{kep} to the outer infalling region using a $v_{\text{rot}}(r) \propto r^{-1}$ rotational profile, but such velocity profiles cannot reproduce the low-velocity redshifted component to the north apparent in the observed $P-V$ diagram along the major axis. This failure to connect the Keplerian to the infalling region using a sensible rotational profile may indicate an abrupt increase in the rotational velocity progressing inward from the infalling region to r_{kep} .

4.2. On the Formation of Keplerian Disks

Although theoretical models have so far only investigated the possible formation of Keplerian circumbinary disks around single protostars, we anticipate that these models will reach the same conclusions on the formation of Keplerian circumbinary disks around the closely separated components of a binary protostellar system such as L1551 NE. Indeed, the mass ($\sim 0.03\text{--}0.12 M_{\odot}$; Paper I) and size (outer radius $\sim 300 \text{ AU}$) of the Keplerian circumbinary disk around L1551 NE, a Class I object, is comparable to the masses and sizes of circumbinary disks observed around single Class I protostars.

As discussed in Section 1, despite the ever increasing evidence for Keplerian circumbinary disks around Class I protostars, many theoretical models continue to have difficulty forming such disks around protostars. Mellon & Li (2009) and Li et al. (2011) have conducted thorough two-dimensional axisymmetric simulations of the collapse of rotating and magnetized singular isothermal cores, in which they included non-ideal MHD effects such as ambipolar diffusion, the Hall effect, and Ohmic dissipation. Their detailed investigations show that even in the case of moderately magnetized cases (mass-to-flux ratio $\lambda \lesssim 10$), magnetic braking is strong enough to remove essentially all of the angular momentum of the material that accretes onto the central object and thus suppress the formation of any sizable (\gtrsim several AU) centrifugally supported disks. They argued that effects not included in their models such as outflow stripping of protostellar envelopes that anchor the magnetic field and/or enhanced magnetic diffusivity due to turbulence-induced reconnection may be required to form Keplerian disks with radii $\gtrsim 100 \text{ AU}$. Machida & Matsumoto (2011) and Machida et al. (2011) performed three-dimensional magnetohydrodynamic simulations of the evolution of Bonnor–Ebert spheres with different initial rotations and magnetic fields. Under a broad range of conditions, they were able to successfully form centrifugally supported disks with radii $\gtrsim 100 \text{ AU}$, although not until near or at the end of the main accretion phase. Because the mass of the circumbinary disk in their model was considerably larger (by a factor of $\sim 2\text{--}100$) than that of the central protostar (see also Vorobyov 2009), however, the circumbinary disk did not exhibit Keplerian motion and was subject to fragmentation. As mentioned above, Li et al. (2011) argued that the simulations performed by Machida & Matsumoto (2011) and Machida et al. (2011) are severely affected by numerical diffusivity of the magnetic fields.

In the case of L1551 NE, our toy model that adequately reproduces the observed $P-V$ diagrams comprises a circumbinary disk having motions dominated by Keplerian rotation inside a radius of ~ 300 AU (the Keplerian radius) and infall (with an infall velocity considerably smaller than the free-fall velocity) outside this radius. Such a break between predominantly rotational and predominantly infalling motions can be seen in the theoretical simulations of Machida & Matsumoto (2011) and Machida et al. (2011), although, as mentioned above, these simulations produce much more massive centrifugally supported circumstellar disks than are actually observed, which moreover do not exhibit Keplerian motion. The abrupt decrease in rotational velocity between the inner and outer regions of this disk indicates that angular momentum is effectively transported outward, presumably by magnetic braking as invoked in theoretical models, beyond the Keplerian radius. Within the Keplerian radius, we cannot easily quantify any deviation in the measured velocity field from Keplerian motion, although some loss of angular momentum is required for material to move through the circumbinary disk and eventually accrete onto the protostellar system.

To explain the difference in the effectiveness at which angular momentum is transported outward within and beyond the Keplerian radius in the circumbinary disk, we borrow ideas from Machida & Matsumoto (2011) and Machida et al. (2011), as well as adding our own. One reason may be the difference in the plasma β between the inner and outer regions of the disk, which is defined as

$$\beta = \frac{8\pi P}{B^2}, \quad (3)$$

where B and P are the magnetic flux density and the thermal pressure, respectively (Machida & Matsumoto 2011; Machida et al. 2011). Strictly speaking, the plasma β applies only to highly-ionized gas. In the present situation of weakly-ionized and primarily molecular gas, the magnetic field is well coupled to the molecular gas through free electrons and ions, at least in directions perpendicular to the magnetic field. Under such circumstances, the gas pressure refers not only to the ionized component, but also to the neutral component; i.e., the total gas pressure, which in the present situation is dominated by the pressure of the molecular gas. To that, we add the likelihood that the ionization fraction decreases dramatically between the outer and inner regions of the disk. At smaller radii, the thermal pressure is likely to be higher because of the higher gas density and temperature; consistent with this idea, the inferred infall velocity is smaller than the expected free-fall velocity near the Keplerian radius, presumably because the infalling material is decelerated as it moves into regions of higher gas pressure. In addition, Ohmic dissipation of magnetic fields becomes increasingly more effective in denser regions closer to the protostar where cosmic rays do not penetrate as easily to ionize the molecular gas (Machida & Matsumoto 2011; Machida et al. 2011). In denser regions, free electrons also become increasingly depleted by adsorption onto dust grains. As a consequence, the plasma β can increase and the ionization fraction can decrease significantly toward smaller radii and hence the control exerted by magnetic fields on the molecular gas becomes dramatically weaker at smaller radii. In this scenario, magnetic braking becomes ineffective within a certain radius (the Keplerian radius), outside of which magnetic braking can operate effectively; this scenario nonetheless requires that the initial infalling matter that forms the Keplerian disk to have retained appreciable angular momentum. The transition

between good and poor magnetic control of the molecular gas may occur quite abruptly not only at a certain radius, but also at a particular stage in the evolution of a protostellar system as the mass and hence density of the circumbinary disk grow. In our toy model, the transition between predominantly infalling to predominantly Keplerian motion in the circumbinary disk is infinitely sharp, a situation which is of course not physical; rather, our toy model suggests that this transition must occur over an annular region that is too narrow to be probed at the angular resolution of our observations.

Finally, we consider the possible future evolution of the Keplerian circumbinary disk in L1551 NE as it continues to accumulate mass from infalling material. On the assumption of local thermodynamic equilibrium, an excitation temperature of $C^{18}O$ (3–2) of 42 K (Moriarty-Schieven et al. 1994), and a $C^{18}O$ abundance of 1.7×10^{-7} (Crapsi et al. 2004), we estimate a mass for the infalling component detected in our observations of $\sim 0.0023 M_{\odot}$. For an infall velocity derived in our model of 0.6 km s^{-1} at a radius of 300 AU, the mass-infall rate is therefore $\sim 9.6 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$ (e.g., see Takakuwa et al. 2007). This value is, incidentally, comparable to the estimated mass-infall rate onto the Keplerian circumstellar disk of L1527 IRS of $\sim 6.6 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$ (Tobin et al. 2012). The envelope around L1551 NE has a mass (in both molecular gas and dust) of $\sim 0.39 M_{\odot}$, as inferred from its dust continuum emission at 0.85 mm observed with the JCMT (Moriarty-Schieven et al. 2006). If this entire envelope drains onto the Keplerian circumbinary disk at the abovementioned mass-infall rate, the disk can continue to accumulate mass for the next $\sim 0.39 M_{\odot} / 9.6 \times 10^{-7} M_{\odot} \text{ yr}^{-1} \sim 4.1 \times 10^5 \text{ yr}$. This timescale is comparable to the typical timescale estimated for the Class I evolutionary stage ($\sim 2.5\text{--}6.7 \times 10^5 \text{ yr}$; Hatchell et al. 2007). Thus, the Keplerian circumbinary disk of L1551 NE may continue to grow in mass until the end of the main accretion phase. Because the total mass of the Keplerian circumbinary disk would then become comparable to the present mass of the protostellar system, however, this disk may become unstable to fragmentation if all the mass that falls onto the Keplerian circumbinary disk remains there. Of course, rather than growing too much in mass and fragmenting, much of the material in the Keplerian circumbinary disk may accrete onto the protostellar components and/or be ejected or entrained in an outflow.

5. SUMMARY

We have conducted follow-up observations of the binary protostellar system L1551 NE in $C^{18}O$ (3–2) with the SMA in its subcompact and compact configurations. We presented results from the new data alone, as well as results obtained by combining the new data with those taken previously with the SMA in its extended configuration (Paper I). Together, the combined data cover spatial scales spanning the range $\sim 140\text{--}2000$ AU. From the resulting maps and $P-V$ diagrams, we found:

1. An elongated feature around the binary protostellar system with an outer dimension of $\sim 1000 \text{ AU} \times 800 \text{ AU}$, as shown in Figure 1. The dimensions of this feature are significantly larger than those of the Keplerian circumbinary disk reported in Paper I of $\sim 600 \text{ AU} \times 300 \text{ AU}$ imaged only with the extended configuration of the SMA. The major axis of this feature matches well with the major axis of the Keplerian circumbinary disk and is approximately perpendicular to the axis of the [Fe II] jets driven by Source

A as well as the major axis of an outflow cavity centered on Source B. This feature therefore traces a circumbinary disk or disk-like structure whose major axis, and presumably also equatorial plane, is perpendicular to outflows from the binary protostellar components.

2. The velocity field of the circumbinary disk-like structure can be decomposed into two distinct components, as seen in the velocity-channel maps of Figures 2–4 and, most clearly, in the corresponding $P-V$ diagrams of Figure 5. One component comprises high-velocity ($>0.5 \text{ km s}^{-1}$) blueshifted and redshifted emission located toward the north and south, respectively, of the binary protostellar system, as seen in Paper I. The other component comprises low-velocity ($<0.5 \text{ km s}^{-1}$) blueshifted and redshifted emission located toward the west and the east, respectively, of the binary protostellar system, seen here for the first time. The high-velocity blueshifted and redshifted components are distinctly separated in space and aligned along the major axis of the circumbinary disk. Their velocity field is consistent with Keplerian motion and so these components are associated with the Keplerian circumbinary disk identified in Paper I. Conversely, the low-velocity components show weak redshifted emission to the north in addition to strong redshifted emission to the south (the latter just like the high-velocity redshifted component) of the binary protostellar system, as well as an east (red) to west (blue) velocity gradient along the minor axis of the circumbinary disk.

We made a toy model comprising a circumbinary disk-like structure that exhibits pure Keplerian motion (for a binary protostellar mass of $0.8 M_{\odot}$, as derived in Paper I) in its inner region and pure infalling motion at a constant velocity in its outer region, which is surrounded by a diffuse envelope. From this toy model, we generated model maps for different Keplerian radii and infall velocities. We then made mock observations of these model maps with the same uv coverage as that attained in our observations with the SMA to generate simulated model maps. From these maps, we derived simulated model $P-V$ diagrams to compare with the observed $P-V$ diagrams. We found that:

3. The inclusion or exclusion of a diffuse envelope, the parameters of which we could only roughly estimate, does not change in an important manner the simulated model $P-V$ diagrams. The best-fit toy model requires motion in the disk to change from (pure) Keplerian to (pure) infall at a radius of $\sim 300 \text{ AU}$, the Keplerian radius. The low-velocity blueshifted and redshifted components discussed above trace primarily the outer infalling portion of the circumbinary disk. The best-fit infall velocity is 0.6 km s^{-1} , which is smaller than the expected free-fall velocity of $\sim 2.2 \text{ km s}^{-1}$ at the Keplerian radius. We tried, but were unable, to connect the motions within and beyond the Keplerian radius with a power-law rotational profile in a manner compatible with the observed $P-V$ diagrams, suggesting that the transition between infall and rotation occurs over a region much narrower than can be studied in our observations.

We considered the implications of our results on the formation of Keplerian circumbinary disks around protostellar systems:

4. The lack of detectable rotation in the infalling component implies that angular momentum is effectively transported outward in the outer region of the circumbinary disk (i.e., beyond the Keplerian radius). On the other hand, angular momentum is not effectively transported outward in

the inner region of the same disk (i.e., within the Keplerian radius). Borrowing in part ideas from Machida & Matsumoto (2011) and Machida et al. (2011), but also adding our own, we suggest that this difference in effectiveness at which angular momentum is transported outward may be caused by the very different plasma β , the ratio of the gas to magnetic pressure, along with the different ionization fractions of the molecular gas in these two regions. Beyond the Keplerian radius, the plasma β is relatively small and the ionization fraction relatively high so that magnetic fields have effective control of the molecular gas, allowing angular momentum to be efficiently transported outward through magnetic braking. Within the Keplerian radius, the plasma β increases whereas the ionization fraction decreases, so that magnetic fields no longer have effective control of the molecular gas and hence magnetic braking becomes ineffective. Consistent with the idea of strong gas pressure near the Keplerian radius, the inferred infall velocity is smaller than the expected free-fall velocity near the Keplerian radius, presumably because the infalling material is decelerated as it moves into regions of higher gas pressure. The transition between good and poor magnetic control of the molecular gas may occur quite abruptly not only at a certain radius, but also at a particular stage in the evolution of a protostellar system as the mass and hence density of the circumbinary disk grow.

5. The inferred mass-accretion rate onto the Keplerian circumbinary disk is $\sim 9.6 \times 10^{-7} M_{\odot} \text{ yr}^{-1}$. Given that the mass in the surrounding dense condensation is $\sim 0.39 M_{\odot}$, there is sufficient material to feed the Keplerian circumbinary disk at the abovementioned mass-accretion rate for the next $\sim 4.1 \times 10^5 \text{ yr}$, which is comparable with the typical timescale estimated for the Class I evolutionary stage. A large fraction of this material, however, must either accumulate onto the central protostellar system and/or be ejected or entrained in an outflow for the circumbinary disk to remain stable throughout its subsequent evolution.

We close by offering a few thoughts on what future work ought to be conducted to better inform models to reproduce the Keplerian circumbinary disk of L1551 NE as well as to test the ideas proposed above for the differences in effectiveness at which angular momentum is transported outward within and beyond the Keplerian radius in the disk. Our observations with the SMA indicate that the transition from predominantly infalling to predominantly Keplerian motion occurs within an annular strip having an angular width of $\sim 2''$ where the observed radial velocity changes by $\sim 0.5 \text{ km s}^{-1}$ (Figure 5). Observations with ALMA should be able to easily probe the kinematics of this transition region and reveal how the infall and rotational velocities change with radius so as to compare with theoretical models (e.g., such as those by Machida & Matsumoto 2011 and Machida et al. 2011). Single-dish mapping of L1551 NE in C^{18}O (3–2), combined with our SMA data, will permit studies of infall and rotation and hence the transport of angular momentum from the outer region of the envelope (at $\gtrsim 1500 \text{ AU}$) to the Keplerian circumbinary disk (see Takakuwa & Kamazaki 2011). Observations of ionic molecules such as HCO^+ and N_2H^+ at high angular resolutions can be used to look for differences in ionization (e.g., Benson et al. 1998; Caselli et al. 1998; Williams et al. 1998) between the Keplerian circumbinary disk and the surrounding infalling region, and hence test our idea that differences in ionization fractions play an important role in how

effectively magnetic fields can exert control over molecular gas in the different regions.

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