

## ON THE VIABILITY OF THE PAH MODEL AS AN EXPLANATION OF THE UNIDENTIFIED INFRARED EMISSION FEATURES

YONG ZHANG (張泳) AND SUN KWOK (郭新)

Department of Physics, The University of Hong Kong, Pokfulam, Hong Kong, China; zhangy96@hku.hk, sunkwok@hku.hk  
Received 2014 September 15; accepted 2014 October 23; published 2014 December 18

### ABSTRACT

Polycyclic aromatic hydrocarbon (PAH) molecules are widely considered the preferred candidate for the carrier of the unidentified infrared emission bands observed in the interstellar medium and circumstellar envelopes. In this paper, we report the results of fitting a variety of non-PAH spectra (silicates, hydrogenated amorphous carbon, coal, and even artificial spectra) using the theoretical infrared spectra of PAHs from the NASA Ames PAH IR Spectroscopic Database. We show that these non-PAH spectra can be well fitted by PAH mixtures. This suggests that a general match between astronomical spectra and those of PAH mixtures does not necessarily provide definitive support for the PAH hypothesis.

*Key words:* infrared: ISM – ISM: general – ISM: lines and bands – ISM: molecules

### 1. INTRODUCTION

The unidentified infrared emission (UIE) bands at 3.3, 6.2, 7.7–7.9, 8.6, 11.3, and 12.7  $\mu\text{m}$  have been detected in a variety of astronomical objects (for a recent review, see Peeters 2013). These UIE bands have long been known to arise from C–H and C–C vibrational modes of aromatic compounds (Duley & Williams 1981) and are therefore also referred to as aromatic infrared bands. A wide variety of hydrocarbon and carbonaceous materials comprising aromatic units have been proposed as possible carriers of the UIE bands. These include hydrogenated amorphous carbon (HAC; Duley & Williams 1983), quenched carbon composites (Sakata et al. 1987), polycyclic aromatic hydrocarbon (PAH) molecules (Léger & Puget 1984; Allamandola et al. 1985, 1989; Puget & Léger 1989), coal (Papoular et al. 1989), and petroleum fractions (Cataldo & Keheyan 2003). Among these models, the PAH hypothesis has been by far the most popular, and the UIE bands are frequently referred to in the literature simply as PAH bands. The case for the PAH hypothesis has been summarized in Tielens (2008).

In spite of the popularity of the PAH hypothesis, there has been increasing doubts about whether PAH molecules can adequately explain the UIE phenomenon. The UIE bands in astronomical spectra are usually accompanied by aliphatic bands at 3.4, 6.9, and 7.3  $\mu\text{m}$  and unassigned features at 15.8, 16.4, 17.4, 17.8, and 18.9  $\mu\text{m}$  (Jourdain de Muizon et al. 1990; Kwok et al. 1999; Chiar et al. 2000; Sturm et al. 2000; Sellgren et al. 2007). In order to account for the aliphatic bands, the PAH model has been modified to include a small number of methyl side groups (Li & Draine 2012; Yang et al. 2013). Furthermore, the UIE bands are often superimposed upon strong broad emission plateaus at 6–9 and 10–15  $\mu\text{m}$  that have been attributed to bending modes emitted by a mixture of alkane and alkene groups (Kwok et al. 2001), implying a significant contribution from the aliphatic component. These observational properties motivate us to invoke an alternative model, mixed aromatic/aliphatic organic nanoparticles (MAONs), as a possible alternative carrier of UIE bands (Kwok & Zhang 2011, 2013). The disorganized molecular structure of MAONs, i.e., aromatic rings ( $sp^2$  bonded) linked by aliphatic chains ( $sp^3$  bonded), are fundamentally different from the chemical structures of PAHs, which are two-dimensional and regular in structure. MAONs represent an

extension of the family of amorphous hydrocarbons that occur naturally in nature and are artificially produced in the laboratory. Amorphous hydrocarbons are often the natural product of combustion, and it is conceivable that similar products can be produced in the circumstellar environment (Colangeli et al. 1997).

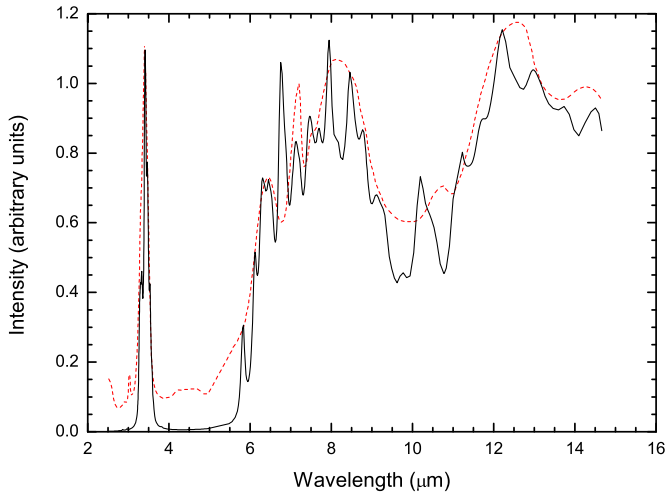
The general agreement between the overall appearance of astronomical spectra and that of PAH mixtures has been widely taken as evidence for the success of the PAH model (see, e.g., Figure 11 of Tielens 2008). Allamandola et al. (1999) and Cami (2011) have reproduced the 5–15  $\mu\text{m}$  spectra of astronomical objects using the experimental and theoretical spectra of PAHs from the NASA Ames PAH IR Spectroscopic Database (PAHdb; Bauschlicher et al. 2010; Boersma et al. 2014). Such a spectral synthesis requires a large variety of PAHs with different sizes, structures, and charge states, each of which has an arbitrary abundance. With a large number of free parameters, such a decomposition suffers from non-uniqueness and cannot be used to identify individual PAH molecules. A more useful exercise is to divide these PAH spectra into a few subclasses according to their sizes, charge states, or compositions and to investigate the variations of these PAH populations in different environments using the PAHdb fitting (Rosenberg et al. 2011, 2014). Boersma et al. (2013) found that the template spectra for the size and charge subclasses created from the PAHdb can be shown to correlate with those derived from astronomical spectra derived from blind signal separation methods (Berné et al. 2007; Pilleri et al. 2012). Specifically, the spatially resolved spectra of the reflection nebula NGC 7023 have been fitted with these templates.

In order to test the validity of the PAH-fitting exercise, we have used the same database and model to fit several spectra obtained from structures that are not made of PAHs. The purpose of the exercise is to test whether the PAHdb model can fit the non-PAH spectra equally well. If it can, then it implies that the fact that the model can fit the astronomical UIE spectra is not evidence for PAH mixtures to be the carrier of UIE.

### 2. TEST OF THE PAHdb MODEL

For the spectral fitting, we made use of version 2.00 of the PAHdb and the IDL package *AmesPAHdbIDLSuite* developed by Boersma et al. (2014).<sup>1</sup> The database contains 700

<sup>1</sup> <http://www.astrochem.org/pahdb>



**Figure 1.** Fitting of the laboratory infrared spectrum (solid line) of HAC obtained by Dischler et al. (1983) using the PAHdb model (dashed line). The value of  $\chi^2/\text{dof}$  is 0.64.

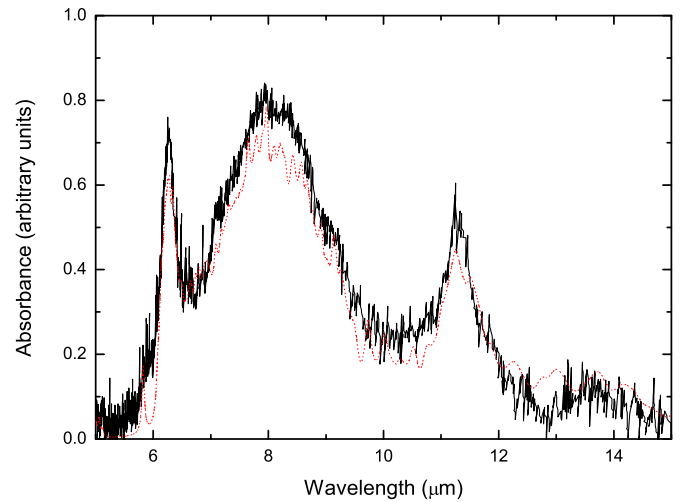
computational and 75 experimental spectra of PAH molecules and ions. The sizes of the PAH molecules range from 6 to 384 carbon atoms, and the charged states include neutral, anion ( $-$ ), and cations ( $+$ ,  $++$ , and  $+++$ ). This is the primary database and tool that have been used to fit astronomical UIE spectra in support of the PAH model (see, e.g., Cami 2011; Boersma et al. 2013).

A non-negative least-squares algorithm was applied to synthesize the target spectra from a total of 700 theoretical PAH spectra in PAHdb. We assume that all PAHs thermally emit at a temperature of 500 K. In order to be consistent with previous fittings using the PAHdb model, we have convolved the PAH spectra by a Lorentzian profile with a fixed width of  $15 \text{ cm}^{-1}$  and redshift all features by  $15 \text{ cm}^{-1}$  to compensate the anharmonic effect (Rosenberg et al. 2014).

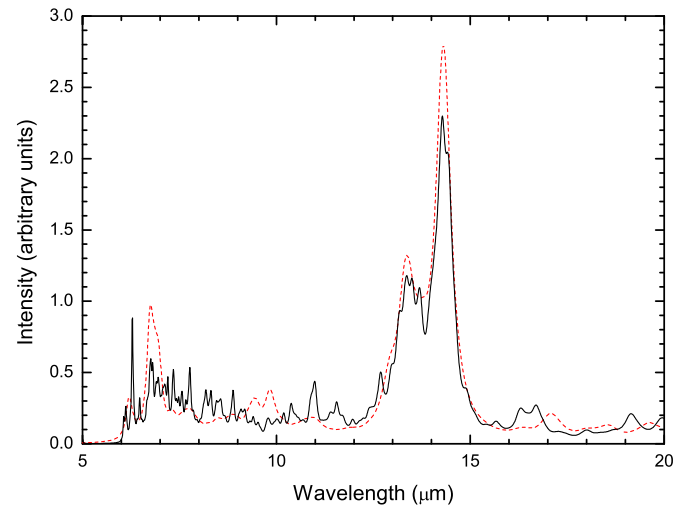
In order to quantitatively evaluate the goodness of the fittings, we computed the  $\chi^2/\text{dof}$  for each of the fittings, where dof is the degree of freedom, defined as the number of wavelength bins of the observed spectra. The values of  $\chi^2/\text{dof}$  in each of the fittings are given in the respective figure captions.

The PAHdb fitting was performed on several non-PAH spectra, including the laboratory spectra of HAC and coal, the theoretical spectrum of a  $sp^2/sp^3$  bonded molecule  $\text{C}_{55}\text{H}_{56}$ , an astronomical spectrum of amorphous silicates, and a number of artificially generated spectra with random wavelengths and intensities. In general, neutral, anionic, and cationic PAHs are all needed in the fitting. Typically, the models include about 40 different molecular species. In all the cases, the contributions from small PAHs ( $N_c < 50$ ) are more significant than large PAHs. This is similar to the fittings of UIE spectra in previous works (Boersma et al. 2014). The relative contributions from different charge states change from case to case.

The HAC spectrum taken from Dischler et al. (1983) shows two strong plateau features at 6–9 and 10–15  $\mu\text{m}$  (Figure 1). Such plateau features are commonly found in the spectra of aliphatic-rich soot samples (Pino et al. 2008), in spectra of coal (Guillois et al. 1996), and in spectra of astronomical sources showing strong aliphatic bands (Kwok et al. 2001). Although the spectra of individual PAH molecules do not show such broad bands, Figure 1 shows that the spectrum of the HAC materials obtained in laboratory can be well fitted using the PAHdb model. Given the large number and diverseness of the



**Figure 2.** Fitting of the absorbance spectra of anthracite coal (solid line; from Guillois et al. 1996) using the PAHdb model (dashed line). The value of  $\chi^2/\text{dof}$  is 0.09.

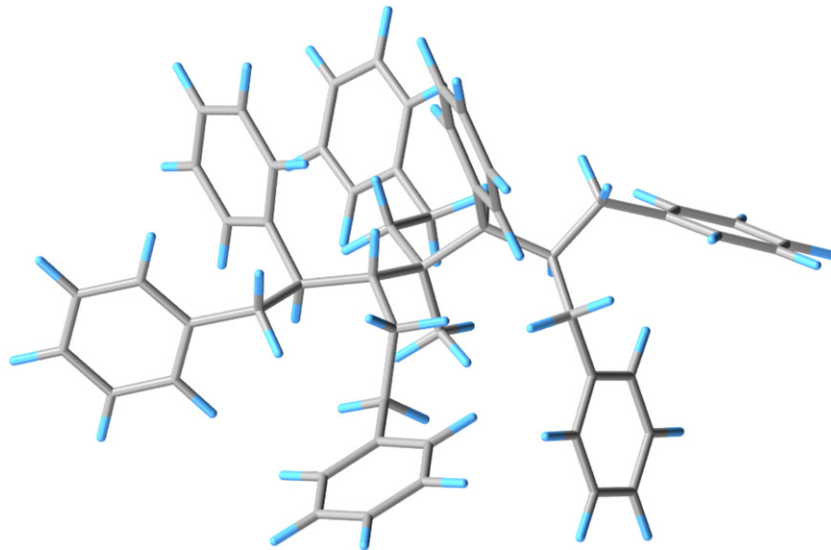


**Figure 3.** Fitting of the spectrum (solid line) of the  $sp^2/sp^3$  bonded molecule  $\text{C}_{55}\text{H}_{56}$  by the PAHdb model (dashed line). The value of  $\chi^2/\text{dof}$  is 0.05.

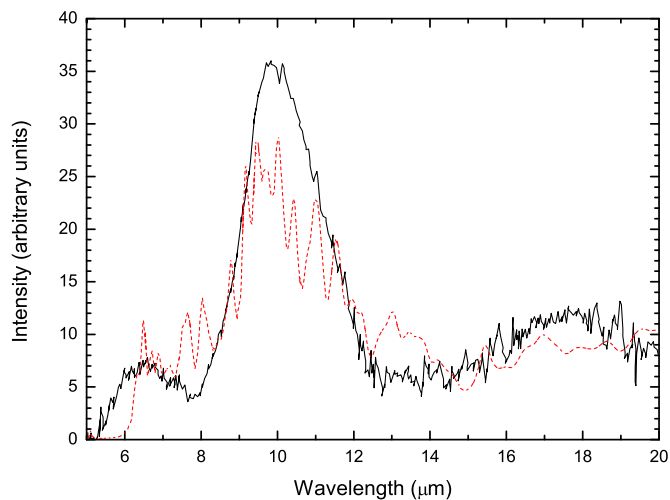
PAH molecules used in the PAHdb model, emission features can be found at almost any wavelength within the 5–15  $\mu\text{m}$  range. The emission plateaus can easily be artificially reproduced by blending many PAH bands together. Therefore, the fitting of infrared emission using PAH spectra may not necessarily imply that PAH molecules are the emitter.

Figure 2 shows the fit to the absorbance spectrum of anthracite coal (Guillois et al. 1996) with the PAHdb model. The spectrum shows some resemblance to the HAC spectrum in having broad plateau features. Again, a mixture of neutral, anion, and cation PAHs can fit the spectrum of coal.

The PAHdb model can also easily reproduce the theoretical spectrum of  $\text{C}_{55}\text{H}_{56}$  (Figure 3).  $\text{C}_{55}\text{H}_{56}$  has a three-dimensional disordered structure with aromatic rings linked by aliphatic chains (Figure 4), representing a possible partial structure of an MAON particle. The chemical structure of this molecule is completely different from the family of PAH molecules. Its theoretical spectrum is obtained from the density functional theory following the method described in Sadjadi et al. (2014). Although the strong aliphatic feature at  $6.9 \mu\text{m}$  cannot be accounted for by PAH mixture, the overall spectrum of  $\text{C}_{55}\text{H}_{56}$  can be well reproduced.



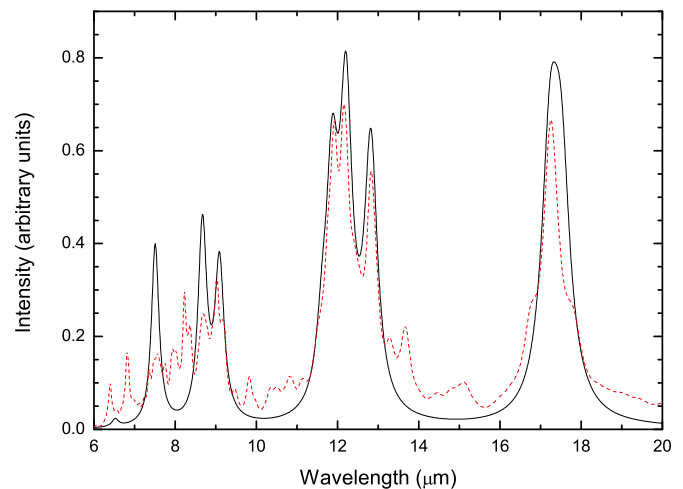
**Figure 4.** Molecular structure of the  $C_{55}H_{56}$  molecule. Carbon and hydrogen atoms are represented in gray and blue, respectively. Figure provided by SeyedAbdolreza Sadjadi.



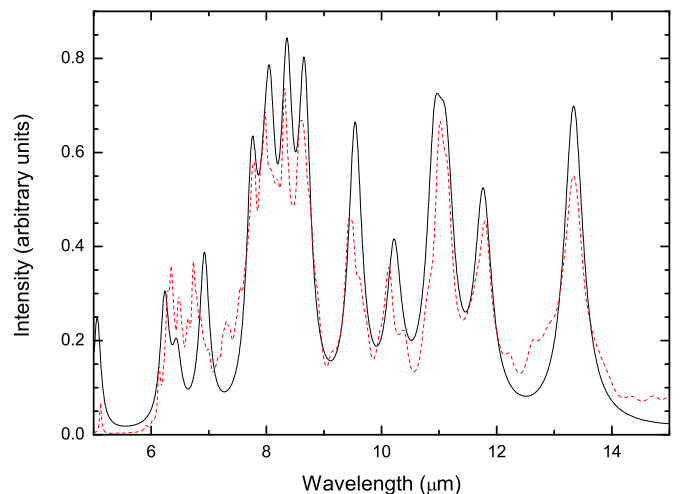
**Figure 5.** PAHdb model fitting (dashed line) of the continuum-subtracted *ISO* spectrum (solid line) of the O-rich AGB star V778 Cyg. The features at 10 and 18  $\mu\text{m}$  are due to amorphous silicates. The value of  $\chi^2/\text{dof}$  is 1.56.

Figure 5 shows the fitting of the emission spectrum from amorphous silicates in the circumstellar envelopes of the asymptotic giant branch (AGB) star V778 Cyg (Yamamura et al. 2000). Silicate features, in either emission or absorption, are commonly observed in oxygen-rich AGB stars (Kwok et al. 1997). The spectrum of V778 Cyg shows both the 10 and 18  $\mu\text{m}$  amorphous silicate features in emission, and the PAHdb model has no problem creating a fit for the spectrum.

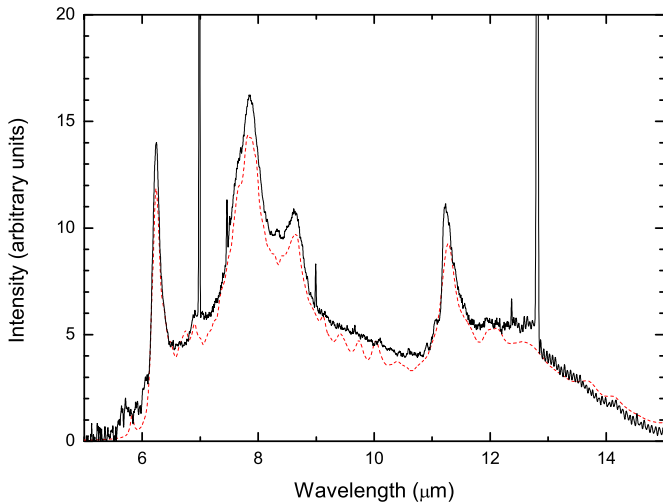
We have also performed a number of exercises where the PAHdb model is used to artificially fit spectra with 10 and 20 randomly generated emission bands. These features are assumed to have a Drude profile with a fixed fractional width of 0.03, and their wavelengths and intensities are completely random within the range of 5–15  $\mu\text{m}$ . Examples of two of such fittings are shown in Figures 6 and 7. Reasonable matches can be achieved between the artificial spectra and the PAH mixture. This suggests that any arbitrary mid-infrared spectra could be successfully simulated using a PAH model through adjusting the weights of the individual PAHs. In the example of Figure 6, the strong



**Figure 6.** PAHdb model fitting (dashed line) of an artificial spectrum (solid line) generated by 10 random features in the 6–20  $\mu\text{m}$  range. The value of  $\chi^2/\text{dof}$  is 0.04.



**Figure 7.** PAHdb model fitting (dashed line) of an artificial spectrum (solid line) generated by 20 random features in the 5–15  $\mu\text{m}$  range. The value of  $\chi^2/\text{dof}$  is 0.11.



**Figure 8.** Continuum-subtracted *ISO* spectrum of BD+30°3639 (black curve) fitted by PAHdb model (red curve). The narrow lines in the spectrum of BD+30°3639 are atomic lines and are excluded from the fitting. The value of  $\chi^2/\text{dof}$  is 0.36.

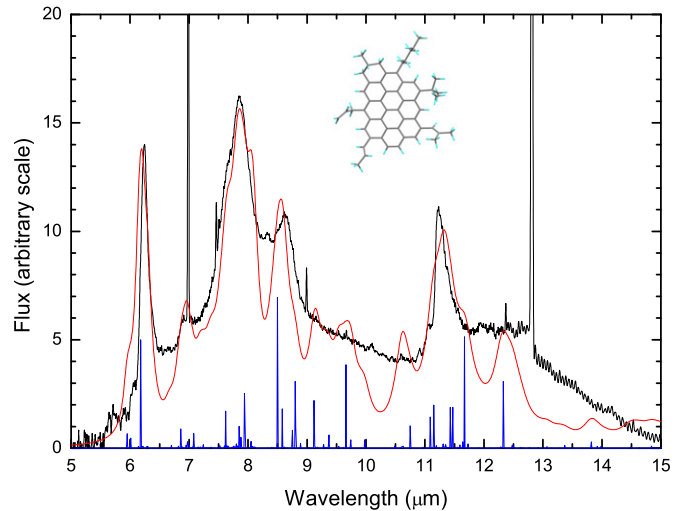
artificial features at 12 and 17  $\mu\text{m}$  clearly have nothing to do with aromatic compounds and yet the PAHdb model can fit them well. In Figure 7, there are multiple strong artificial features; however, the PAHdb model is still flexible enough to produce a satisfactory fit.

Although the PAHdb fitting model can produce reasonably good fits to non-PAH spectra as seen in the above examples, it may be argued that the PAHdb model may produce better fits to UIE spectra. In order to make a quantitative comparison, we also performed a fitting to the continuum-subtracted *Infrared Space Observatory (ISO)* spectrum of the planetary nebula BD+30°3639 using the PAHdb model (Figure 8). The continuum component was subtracted using a linear fit to the line-free regions. The  $\chi^2/\text{dof}$  value of the fit is 0.36, which is comparable or larger than those for non-PAH spectrum fitting. This suggests that the PAH model does not necessarily generate a better match with astronomical spectra than with non-PAH spectra.

### 3. FIT OF ASTRONOMICAL SPECTRUM BY A SIMPLE MAON MOLECULE

In this section, we investigate the possibility of fitting the astronomical spectra of BD+30°3639 with a MAON model. For the model, we chose the theoretical spectrum of  $\text{C}_{55}\text{H}_{52}$ , which is one of a series of molecules with mixed aliphatic/aromatic structures calculated by Sadjadi et al. (2014). After applying a Drude profile of 500 K, the spectral regions that contain high-density and high-intensity modes will have strong bands (Figure 9).

The theoretical spectrum of the molecule shows two broad plateau features centered around 7 and 11  $\mu\text{m}$  and can be made to fit the observed spectrum by redshifting features with wavelengths larger than 10  $\mu\text{m}$  by 0.4  $\mu\text{m}$ . The strengths of some of the features are also artificially adjusted: the 6.2  $\mu\text{m}$  feature and the features between 7.5 and 9.4  $\mu\text{m}$  by a factor of 2.3 and the 6.9  $\mu\text{m}$  feature by a factor of 0.5 (Figure 9). Such shifting and adjustment could be the result of anharmonic effect, impurity, geometry, and/or changes in the  $sp^2/sp^3$  ratio. Since  $\text{C}_{55}\text{H}_{52}$  is a simple molecule and is only one of many possible MAON structures, we are not claiming that the carrier of UIE bands is



**Figure 9.** Continuum-subtracted *ISO* spectrum of BD+30°3639 (black curve) fitted by a modified spectrum of  $\text{C}_{55}\text{H}_{52}$  (red curve). The chemical structure of the molecule is shown in the insert. The relative intensities of theoretical vibrational modes are marked by vertical blue solid lines. The curve of  $\text{C}_{55}\text{H}_{52}$  is obtained by applying a Drude profile of 500 K to the vibrational lines (for details, see Sadjadi et al. 2014). The narrow lines in the spectrum of BD+30°3639 are atomic lines and are excluded from the fitting.

this specific molecule. What this example shows is that instead of involving hundreds of molecules and using an equally large number of free parameters, astronomical UIE spectra can be fitted with a single molecule with a few parameters. This could be a simple coincidence, or the MAON structure has intrinsic properties that are similar to those of the actual carrier of the UIE bands. This exercise illustrates that there are other alternatives to fit the UIE bands than relying on mixtures of PAH molecules.

## 4. DISCUSSION

The fact that the PAHdb model can produce fits to HAC, coal, molecules of mixed  $sp^2/sp^3$  structure, silicates, and even randomly generated spectra suggests that the model contains far too many parameters to be considered useful. One of the objections of Rosenberg et al. (2014) to the MAON model is that it relies on an empirical decomposition of the observed spectra. Here, we show that the spectral decomposition based on theoretical PAH spectra does not produce more reliable results than the empirical decomposition. Through these examinations, we can conclude that a match between astronomical spectra and a PAH mixture cannot be taken as an evidence of PAHs as the carrier of UIE.

It is generally acknowledged that no single PAH can reproduce the UIE spectra and a collection of PAH molecules of various sizes, charge states, and compositions is required to fit astronomical spectra. So far, there has been no rigorous study on whether the diversity of PAHs is compatible with different astronomical environments. Although the UIE bands are found in highly diverse radiation environments, neutral and positive/negative PAH ions are always required to fit the observed spectra.

To date, there is no experimental support for the synthesis of large groups of diverse PAH molecules under natural conditions. In contrast, amorphous carbonaceous compounds occur naturally as the result of combustion or as the result of injection of energy (arc discharge, laser ablation, laser pyrolysis, photolysis, or microwave) into mixed hydrocarbon gas-phase molecules



or graphite (Sakata et al. 1987; Scott & Duley 1996; Herlin et al. 1998; Mennella et al. 1999, 2003; Jäger et al. 2009; Dartois et al. 2004; Pino et al. 2008).

The PAH model has particular difficulty in explaining the plateau features. The plateau features are often explained as the blending of PAH modes, and Lorentzian-like broad wings have to be artificially assumed (Li & Draine 2012). In the MAON model, the plateau features arise naturally from the superposition of many in-plane and out-of-plane bending modes of aliphatic groups as these modes are clustered around 8 and 12  $\mu\text{m}$  regions (Kwok et al. 2001). This is supported by the fact that the UIE and plateau features can be fitted by spectral components decomposed from laboratory spectra of HAC samples (Duley & Hu 2012).

If the UIE features are due to PAH molecules, then PAH molecules must be prevalent in the interstellar medium (ISM) and should be easily detectable in absorption against bright stars. The observed strengths of the UIE bands would require a PAH to  $\text{H}_2$  abundance ratio of  $\sim 3 \times 10^{-7}$  (Tielens 2008). However, no individual PAH molecule has been detected through their electronic, vibrational, or rotational transitions (Clayton et al. 2003; Salama et al. 2011; Gredel et al. 2011). These failures have been explained by the proposition that the UIE bands are the result of a collection of a large number of diverse PAH molecules, and the abundance of individual PAH molecules is too low to be detected. Simulations also show that it is impossible to identify individual PAH molecules through their collective spectra in the fitting process of astronomical UIE spectra (Rosenberg et al. 2014).

While PAH molecules have narrow vibrational modes occurring at different frequencies, MAONs, in spite of having different chemical compositions, have naturally broad and similar spectral behavior. Through the introduction of aliphatic groups, the modes around 8 and 12  $\mu\text{m}$  are greatly enriched due to blended vibrational modes of aliphatic structure and activation of Raman modes. While a collection of nine small PAHs is not able to reproduce the observed spectrum of BD+30° 3639 (Cook & Saykally 1998), a simple mixed  $sp^2/sp^3$  molecule  $\text{C}_{55}\text{H}_{52}$  can (Figure 9). While this in itself may not mean anything, it does suggest that it would be worthwhile to explore the spectral properties of MAONs. Since MAONs represent a family of amorphous structures, we need to further study how the spectra would vary in response to changes in size, aliphatic/aromatic ratio, and geometry.

## 5. CONCLUSION

Among the candidate carriers of UIE, the PAH model has attracted the most attention, in part because of its aesthetic appeal. PAH molecules are small and simple gas-phase molecules, and their presence in space is easier to accept than complex organics. However, we now know that complex organics of abiological origin are prevalent in the solar system. Because the PAH hypothesis does not specify any single molecule but refers to a large number of diverse molecules, it does not produce concrete predictions for testing or falsification. Although a mixture of many PAHs can reproduce the UIE spectra, we have shown that such a mixture is also able to produce non-PAH spectra, even randomly generated spectra. Thus, the observational foundation of the PAH model is not solid.

Since the UIE phenomenon is seen throughout the universe, even during its early epochs, a correct identification of the carrier is of great importance. Due to the strengths of the UIE bands, the carrier represents a major reservoir of carbon. Whether

the carriers are a collection of free-flying gas-phase PAH molecules or complex organic solid particles has significant different implications on our understanding of cosmic chemical synthesis, energy exchange between stars and the ISM, and galactic chemical enrichment. The PAH hypothesis has been entrenched in models of the ISM that PAH molecules are believed to be dominant factors in the photoelectric heating of interstellar gas and a determining factor in the ionization balance inside molecular clouds (Joblin & Tielens 2011). In this paper, we try to show that the situation is not as simple as it has been represented in the literature. A correct interpretation of the UIE spectrum requires further studies, both computationally and experimentally, of the vibrational properties of amorphous carbons of mixed aromatic/aliphatic structures.

We thank SeyedAbdolreza Sadjadi and Renaud Papoular for helpful discussions. The fitting models were done using the NASA Ames Research Center PAH IR Spectroscopic Database and software, and we thank Christiaan Boersma for making these resources publicly available. This work was partially supported by the Research Grants Council of the Hong Kong Special Administrative Region, China (project No. HKU 7031/10P).

## REFERENCES

- Allamandola, L. J., Hudgins, D. M., & Sandford, S. A. 1999, *ApJL*, 511, L115  
 Allamandola, L. J., Tielens, A. G. G. M., & Barker, J. R. 1985, *ApJL*, 290, L25  
 Allamandola, L. J., Tielens, A. G. G. M., & Barker, J. R. 1989, *ApJS*, 71, 733  
 Bauschlicher, C. W., Jr., Boersma, C., Ricca, A., et al. 2010, *ApJS*, 189, 341  
 Berné, O., Joblin, C., Deville, Y., et al. 2007, *A&A*, 469, 575  
 Boersma, C., Bauschlicher, C. W., Jr., Ricca, A., et al. 2014, *ApJS*, 211, 8  
 Boersma, C., Bregman, J. D., & Allamandola, L. J. 2013, *ApJ*, 769, 117  
 Cami, J. 2011, in *EAS Publication Ser.*, Vol. 46, PAHs and the Universe, ed. C. Joblin & A. G. G. M. Tielens (Cambridge: Cambridge Univ. Press), 117  
 Cataldo, F., & Keheyan, Y. 2003, *IJASB*, 2, 41  
 Chiar, J. E., Tielens, A. G. G. M., Whittet, D. C. B., et al. 2000, *ApJ*, 537, 749  
 Clayton, G. C., Gordon, K. D., Salama, F., et al. 2003, *ApJ*, 592, 947  
 Colangeli, L., Bussoletti, E., Pestellini, C. C., et al. 1997, *AdSpR*, 20, 1617  
 Cook, D. J., & Saykally, R. J. 1998, *ApJ*, 493, 793  
 Dartois, E., Muñoz Caro, G. M., Deboffle, D., & d'Hendecourt, L. 2004, *A&A*, 423, L33  
 Dischler, B., Bubenzer, A., & Koidl, P. 1983, *SSCom*, 48, 105  
 Duley, W. W., & Hu, A. 2012, *ApJ*, 761, 115  
 Duley, W. W., & Williams, D. A. 1981, *MNRAS*, 196, 269  
 Duley, W. W., & Williams, D. A. 1983, *MNRAS*, 205, 67  
 Gredel, R., Carpentier, Y., Rouillé, G., et al. 2011, *A&A*, 530, 26  
 Guillois, O., Nenner, I., Papoular, R., & Reynaud, C. 1996, *ApJ*, 464, 810  
 Herlin, N., Bohn, I., Reynaud, C., et al. 1998, *A&A*, 330, 1127  
 Jäger, C., Huisken, F., Mutschke, H., Jansa, I. L., & Henning, T. H. 2009, *ApJ*, 696, 706  
 Joblin, C., & Tielens, A. G. G. M. 2011, PAHs and the Universe (Les Ulis, France: EDP Sciences)  
 Jourdain de Muizon, M., D'Hendecourt, L. B., & Geballe, T. R. 1990, *A&A*, 235, 367  
 Kwok, S., Volk, K., & Bernath, P. 2001, *ApJL*, 554, L87  
 Kwok, S., Volk, K., & Bidelman, W. P. 1997, *ApJS*, 112, 557  
 Kwok, S., Volk, K., & Hrivnak, B. J. 1999, *A&A*, 350, L35  
 Kwok, S., & Zhang, Y. 2011, *Natur*, 479, 80  
 Kwok, S., & Zhang, Y. 2013, *ApJ*, 771, 5  
 Léger, A., & Puget, J. L. 1984, *A&A*, 137, L5  
 Li, A., & Draine, B. T. 2012, *ApJL*, 760, L35  
 Mennella, V., Baratta, G. A., Esposito, A., Ferini, G., & Pendleton, Y. J. 2003, *ApJ*, 587, 727  
 Mennella, V., Brucato, J. R., Colangeli, L., & Palumbo, P. 1999, *ApJL*, 524, L71  
 Papoular, R., Conrad, J., Giuliano, M., Kister, J., & Mille, G. 1989, *A&A*, 217, 204  
 Peeters, E. 2013, in *IAU Symp. 297, The Diffuse Interstellar Bands*, ed. J. Cami & N. L. J. Cox (Cambridge: Cambridge Univ. Press), 187  
 Pilleri, P., Montillaud, J., Berné, O., & Joblin, C. 2012, *A&A*, 542, 69  
 Pino, T., Dartois, E., Cao, A.-T., et al. 2008, *A&A*, 490, 665

- Puget, J. L., & Léger, A. 1989, [ARA&A](#), **27**, 161
- Rosenberg, M. J. F., Berné, O., & Boersma, C. 2014, [A&A](#), **566**, L4
- Rosenberg, M. J. F., Berné, O., Boersma, C., Allamandola, L. J., & Tielens, A. G. G. M. 2011, [A&A](#), **532**, 128
- Sadjadi, S., Zhang, Y., & Kwok, S. 2014, *ApJ*, submitted
- Sakata, A., Wada, S., Onaka, T., & Tokunaga, A. T. 1987, [ApJL](#), **320**, L63
- Salama, F., Galazutdinov, G. A., Krelowski, J., et al. 2011, [ApJ](#), **728**, 154
- Scott, A., & Duley, W. W. 1996, [ApJL](#), **472**, L123
- Sellgren, K., Uchida, K. I., & Werner, M. W. 2007, [ApJ](#), **659**, 1338
- Sturm, E., Lutz, D., Tran, D., et al. 2000, *A&A*, **358**, 481
- Tielens, A. G. G. M. 2008, [ARA&A](#), **46**, 289
- Yamamura, I., Dominik, C., de Jong, T., Waters, L. B. F. M., & Molster, F. J. 2000, *A&A*, **363**, 629
- Yang, X. J., Glaser, R., Li, A., & Zhong, J. X. 2013, [ApJ](#), **776**, 110